



The impact of baryonic physics on the structure of dark matter haloes: the view from the FIRE cosmological simulations

T. K. Chan,^{1★} D. Kereš,^{1★} J. Oñorbe,² P. F. Hopkins,³ A. L. Muratov,¹
C.-A. Faucher-Giguère⁴ and E. Quataert⁵

¹Department of Physics, Center for Astrophysics and Space Sciences, University of California at San Diego, 9500 Gilman Drive, La Jolla, CA 92093, USA

²Max-Planck-Institut fuer Astronomie, Koenigstuhl 17, D-69117 Heidelberg, Germany

³TAPIR, Mailcode 350-17, California Institute of Technology, Pasadena, CA 91125, USA

⁴Department of Physics and Astronomy and CIERA, Northwestern University, 2145 Sheridan Road, Evanston, IL 60208, USA

⁵Department of Astronomy and Theoretical Astrophysics Center, University of California Berkeley, Berkeley, CA 94720, USA

Accepted 2015 September 16. Received 2015 September 14; in original form 2015 July 9

ABSTRACT

We study the distribution of cold dark matter (CDM) in cosmological simulations from the FIRE (Feedback In Realistic Environments) project, for $M_* \sim 10^{4-11} M_\odot$ galaxies in $M_h \sim 10^{9-12} M_\odot$ haloes. FIRE incorporates explicit stellar feedback in the multiphase interstellar medium, with energetics from stellar population models. We find that stellar feedback, without ‘fine-tuned’ parameters, greatly alleviates small-scale problems in CDM. Feedback causes bursts of star formation and outflows, altering the DM distribution. As a result, the inner slope of the DM halo profile (α) shows a strong mass dependence: profiles are shallow at $M_h \sim 10^{10-10^{11}} M_\odot$ and steepen at higher/lower masses. The resulting core sizes and slopes are consistent with observations. This is broadly consistent with previous work using simpler feedback schemes, but we find steeper mass dependence of α , and relatively late growth of cores. Because the star formation efficiency M_*/M_h is strongly halo mass dependent, a rapid change in α occurs around $M_h \sim 10^{10} M_\odot$ ($M_* \sim 10^6-10^7 M_\odot$), as sufficient feedback energy becomes available to perturb the DM. Large cores are not established during the period of rapid growth of haloes because of ongoing DM mass accumulation. Instead, cores require several bursts of star formation after the rapid build-up has completed. Stellar feedback dramatically reduces circular velocities in the inner kpc of massive dwarfs; this could be sufficient to explain the ‘Too Big To Fail’ problem without invoking non-standard DM. Finally, feedback and baryonic contraction in Milky Way-mass haloes produce DM profiles slightly shallower than the Navarro–Frenk–White profile, consistent with the normalization of the observed Tully–Fisher relation.

Key words: galaxies: evolution – galaxies: haloes – galaxies: kinematics and dynamics – galaxies: structure – dark matter.

1 INTRODUCTION

Cold dark matter with a cosmological constant (Λ CDM) is a successful cosmological model that can simultaneously explain large-scale fluctuations in the cosmic microwave background and the large-scale structure of the Universe that forms out of these fluctuations at much later time (Springel et al. 2005; Spergel et al. 2007). However, on much smaller scales, within DM haloes that host observed galaxies, there are indications that the distribution of DM is inconsistent with the simplest prediction of the CDM paradigm.

The most obvious and most studied disagreement is in density profiles of DM haloes inferred from observations of dwarf and low surface-brightness galaxies. While observed slopes are relatively flat (central density slope $\alpha \sim 0$, where central density $\propto r^\alpha$; Salucci & Burkert 2000; Swaters et al. 2003; Gentile et al. 2004; Spekkens, Giovanelli & Haynes 2005; de Blok et al. 2008; Walter et al. 2008; Oh et al. 2011), simulated CDM haloes are cuspy ($\alpha \sim -1$; Flores & Primack 1994; Moore 1994; Navarro, Frenk & White 1997). This problem is known as the *cusp/core* problem.

To address this problem, various modifications of DM properties have been proposed to erase the steep central regions and produce a core-like density profile. Examples include warm dark matter (WDM), whose free streaming can suppress the small-scale

*E-mail: tkc004@physics.ucsd.edu (TKC); dkeres@physics.ucsd.edu (DK)

structure (Abazajian 2006; Dunstan et al. 2011; Lovell et al. 2012) and self-interacting dark matter (SIDM) whose interaction can substantially affect the central density profile of haloes (e.g. Burkert 2000; Kochanek & White 2000; Yoshida et al. 2000; Davé et al. 2001; Elbert et al. 2015). There is still no consensus on whether these modifications can solve the problem and satisfy all observational constraints: these modifications can produce serious problems on their own. For example Macciò et al. (2012a) found that in order to produce DM cores as large as those seen in observed dwarf galaxies, the WDM would also prevent the formation of dwarf galaxies. Simple SIDM models (Carlson, Machacek & Hall 1992; Machacek, Carlson & Hall 1993; de Laix, Scherrer & Schaefer 1995) were shown to violate observations of central regions of galaxy clusters (Yoshida et al. 2000; Miralda-Escudé 2002) that are found to be denser and more elliptical than SIDM would predict. However, recent SIDM models that take into account more accurate observational constraints and the effects of baryons offer promising explanation of the problem without violating any known observational constraints (Peter et al. 2013; Rocha et al. 2013; Kaplinghat et al. 2014; Elbert et al. 2015).

Initial problems with the SIDM have motivated more complex models such as velocity-dependent SIDM (Yoshida et al. 2000; Loeb & Weiner 2011; Macciò et al. 2012a). However, SIDM with a simple power-law velocity dependence will not be able to create a core and, at the same time, produce stable haloes of dwarf galaxies over a Hubble time (Gnedin & Ostriker 2001). Loeb & Weiner (2011) proposed SIDM with a Yukawa potential, which has a non-trivial velocity dependence that is effective at producing cores in dwarf galaxies without adverse effects on clusters of galaxies. Cosmological simulation of a Milky Way-mass halo with this SIDM showed that realistic cores can be formed in subhaloes expected to host dwarf galaxies (Vogelsberger, Zavala & Loeb 2012a). However, this model requires more free parameters for the velocity dependence and it is not yet known whether it can reproduce the correct halo abundance and mass distribution.

Before concluding that simple CDM models must be modified, we must also examine the effects of baryons on the distribution of DM within haloes. Baryons are not only what is actually observed in galaxies, but baryonic effects at the halo centre can, in principle, also play a role in shaping the DM profiles (Blumenthal et al. 1986; Navarro, Eke & Frenk 1996; El-Zant, Shlosman & Hoffman 2001; Gnedin et al. 2004; Read & Gilmore 2005; Governato et al. 2010, 2012; Macciò et al. 2012b; Peñarrubia et al. 2012; Pontzen & Governato 2012, 2014; Teyssier et al. 2013; Di Cintio et al. 2014).

Navarro et al. (1996) and Read & Gilmore (2005) used N -body simulations to model a sudden removal of a large baryonic component via supernova (SN)-driven winds (represented as a change in the external potential) in a dwarf galaxy with an initially cuspy DM halo. They showed that such mass removal leads to formation of a DM core. An alternative mechanism was proposed by Mashchenko, Couchman & Wadsley (2006), who showed that bulk motion of gas within forming galaxies leads to significant gravitational potential changes which can also redistribute DM and reduce its central density. Dynamical effects between baryons and DM during the halo formation were also suggested as a mechanism that could modify DM density profiles (e.g. El-Zant, Shlosman & Hoffman 2001; Tonini, Lapi & Salucci 2006; Romano-Díaz et al. 2008; Del Popolo 2009).

Recently Governato et al. (2010, 2012) used cosmological ‘zoom-in simulations’ with baryons, cooling, star formation (SF) and SNe feedback to show that outflow episodes in dwarf galaxies can turn the central DM cusps into cores. Strong SN-driven outflows from

clustered SF in the inhomogeneous interstellar medium (ISM) resulted in a decrease of the DM density within central kiloparsec to less than half of what it would otherwise be in a halo of this mass ($m_{\text{dm}} \approx 10^{10} M_{\odot}$). Pontzen & Governato (2012) further clarified this density flattening mechanism: a quick change in gravitational potential due to gas outflow can effectively inject energy into DM orbits and (typically after many outflow episodes) flatten the central DM profile. They showed that the repeated changes of gravitational potential on time-scales shorter than t_{dyn} during $2 < z < 4$ can significantly flatten cuspy DM profiles.

Brook et al. (2012) showed that a large fraction of the gas that is expelled returns via a large-scale galactic fountain (see also Oppenheimer et al. 2010) to form stars at later times: this greatly increases the chance of outflows from the inner regions and further helps in core formation. Other, non-cosmological simulations with strong SNe feedback also showed that it is possible to form cores in dwarf galaxies owing to bursty SF that removes large quantities of gas during bursts (Teyssier et al. 2013).

On the other hand, idealized simulations by Gnedin & Zhao (2002), Ogiya & Mori (2011) and Garrison-Kimmel et al. (2013) focused on time evolution, SNe energy requirements and mass ejection frequency in idealized models and argued that SNe-driven feedback is not efficient enough to form cores at the observed level. However, recent cosmological simulations have had more success. Madau, Shen & Governato (2014) suggest that earlier mass removal can lower the energy requirements for core formation and Oñorbe et al. (2015) showed that late SF, after the early epoch of cusp building, is particularly efficient at utilizing stellar feedback to remove DM.

There are two other related ‘problems’ with structural properties of CDM haloes. One is the lack of very steep central profiles in relatively massive disc galaxies. Cuspy Navarro–Frenk–White (NFW) profiles (Navarro et al. 1997) are expected to be even steeper within baryon-dominated galaxies owing to the contraction of DM caused by the central concentration of baryons (Blumenthal et al. 1986). The distribution of matter in galaxies effectively determines the Tully–Fisher relation (TFR; e.g. Tully & Fisher 1976; Dutton et al. 2007). However, given the observed distribution of baryons, contraction of an NFW halo would result in circular velocities too high at a given luminosity/mass. This motivated several authors to suggest that DM does not undergo contraction or is perhaps even expanded from the original cuspy NFW profile (Dutton et al. 2007). While this could be interpreted as suggesting a problem with the currently favoured CDM model, stellar feedback is also able to effectively ‘expand’ the DM distribution even in Milky Way-mass haloes (Macciò et al. 2012b).

The second problem is the so-called Too Big To Fail (TBTf) problem (Boylan-Kolchin, Bullock & Kaplinghat 2011). In the Milky Way, satellites have significantly lower DM densities in the inner few hundred parsecs than the corresponding subhaloes in CDM-only simulations without baryons. Alternatively, massive subhaloes whose inner densities are high, never formed galaxies. Similar problems also exist in other dwarf galaxies in the Local Group (Garrison-Kimmel et al. 2014; Papastergis et al. 2015). There are hints that feedback can help solve this issue along with the cusp/core problem (Brooks & Zolotov 2014; Madau et al. 2014; Oñorbe et al. 2015), although proper statistics are still lacking.

It is becoming clear that the bursty nature of stellar feedback in galaxies can modify the inner regions of DM haloes. However, in general, most simulations used to study this problem so far used crude and often unphysical implementations of stellar feedback. One might worry that this could impact the effect of stellar

feedback. Most of the cosmological simulations used to address the cusp/core issue simply turn off cooling from the gas heated by SNe ejecta until such gas escapes galaxies (Governato et al. 2010, 2012; Macciò et al. 2012b; Pontzen & Governato 2012; Di Cintio et al. 2014). The delayed cooling is unphysically long and results from a misinterpretation of the standard SN remnant results (Martizzi, Faucher-Giguère & Quataert 2015). In addition, most simulations include only SNe feedback while other stellar feedback mechanisms are ignored or implemented crudely: e.g. radiation pressure, cosmic rays and photo-heating are often approximated with pure thermal energy input and additional freely adjustable parameters (Macciò et al. 2012b; Di Cintio et al. 2014).

Furthermore, the particle mass resolution used in some previous studies was insufficient to properly resolve the observed core sizes. Low resolution may hinder the investigation of central density profiles on small scales owing to two-body relaxation effects (Power et al. 2003). In Appendix A, we show the relation between the convergence radius and particle mass that should be used to estimate resolved scales in different simulations.

To study the *cusp/core* problem in a complex cosmological environment we use simulations from the Feedback in Realistic Environments (FIRE) project (Hopkins et al. 2014, H14)¹, supplemented by four new dwarf galaxy simulations. Our simulations use physically motivated stellar feedback, in which energy and momentum input are based on stellar population synthesis models alone, with no adjusted parameters. In addition to SNe energy and momentum we include radiation pressure, stellar winds, photoionization and photo-electric heating processes. In H14 we show that the M_\star – M_h relation in FIRE is in reasonable agreement with observations, for galaxies residing in halo masses $M \lesssim 10^{12} M_\odot$. This result is sensitive to the feedback physics: simulations with SNe alone fail to reproduce the correct relation, unless additional feedback processes are also incorporated. Overall, stellar feedback in FIRE simulations results in bursty SF histories followed by strong outflow episodes (Muratov et al. 2015) that can affect matter distribution within galaxies.

Our simulations are among the highest resolution cosmological ‘zoom-in’ simulations to date evolved down to $z = 0$ with full baryonic physics. In addition to the advantages in implemented physics and resolution, we adopt the P-SPH ‘pressure–entropy’ formulation of smoothed particle hydrodynamics (SPH; Hopkins 2013), which includes a large number of numerical improvements relative to previous SPH studies, which together significantly improve the treatment of cooling and multiphase fluid mixing, and reduce the well-known discrepancies between SPH and grid-based codes.

In this paper we study haloes with masses $10^9 < M_h / M_\odot < 10^{12}$ with full feedback and their DM only analogues, which enables us to directly compare their DM distributions. We find results in broad agreement with previous work, but with some important differences. We find that stellar feedback affects all of the systems we study but large cores develop only in the halo mass range of $\sim 10^{10} M_\odot$ to a few times $10^{11} M_\odot$. Furthermore we show that cores change over time, and that progenitors of massive galaxies once had more prominent cores. We demonstrate how bursty SF and related feedback correlate with changes in DM haloes and show that feedback effectively cancels the effects of adiabatic contraction. Finally we discuss consequences of our results for the cusp/core issue, the TFR, the ‘TBTF’ problem and indirect DM detection.

We find encouraging trends that have the potential to solve most of the apparent small-scale problems of the CDM paradigm.

The paper is organized as follows: Section 2 includes a brief description of the code and implemented stellar feedback as well as the set-up of the simulations. In Section 3, we show the DM density profiles and their time evolution. In Section 4 we study the effects of stellar feedback on the expected contraction of DM and the TFR. In Section 5, we compare our results with previous work, discuss the implications and propose directions for further investigation.

2 SIMULATIONS

2.1 Simulation code

The simulations in this work were run with the newly developed GIZMO code (Hopkins 2015) in a fully conservative, pressure–entropy-based smoothed particle hydrodynamic (P-SPH) mode (Hopkins 2013). P-SPH eliminates artificial surface tension at contact discontinuities that affects traditional density based SPH (Agertz et al. 2007; Sijacki et al. 2012; Saitoh & Makino 2013). We use the artificial viscosity algorithm with a switch from Cullen & Dehnen (2010) which reduces viscosity to close to zero away from shocks and enables accurate shock capturing. The same higher order dissipation switch is used to trigger entropy mixing at the kernel scale following Price (2008). Time-stepping is controlled by the limiter from Durier & Dalla Vecchia (2012), which limits the difference in time-steps between neighbouring particles, further reducing numerical errors. The gravity solver of the GIZMO code is an updated version of the PM+TREE algorithm from GADGET-3 (Springel 2005) and uses fully conservative adaptive gravitational softening for gas (Price & Monaghan 2007). GIZMO’s softening kernel represents the exact solution of the particle mass distributed over the SPH smoothing kernel (Barnes 2012).

The code performs well on standard strong shock, Kelvin–Helmholtz and Rayleigh–Taylor instabilities, and subsonic turbulence tests (for more details see Hopkins 2013). In cosmological ‘zoom-in’ simulations of a Milky Way-size halo without outflows, the code eliminates most of the artificial fragmentation of halo gas seen in traditional SPH simulations (Kaufmann et al. 2006; Sommer-Larsen 2006; Kereš & Hernquist 2009) and increases cooling from the hot halo gas at late times (e.g. Kereš et al. 2005), when compared to the classical SPH (Kereš et al., in preparation). Overall, the resulting halo gas properties are similar to the results from adaptive-mesh and moving-mesh simulations (Agertz, Teyssier & Moore 2009; Kereš et al. 2012; Vogelsberger et al. 2012b; Nelson et al. 2013).

2.2 Baryonic physics

Our simulations incorporate cooling, SF and physical stellar feedback processes that are observed to be relevant in the ISM. Here we briefly review these components; for detailed description please see H14.

Gas follows an ionized+atomic+molecular cooling curve from 10 – 10^{10} K, including metallicity dependent fine-structure and molecular cooling at low temperatures, and at high temperatures ($>10^4$ K) metal-line cooling followed species-by-species for 11 separately tracked species. At all times, we tabulate the appropriate ionization states and cooling rates from a compilation of CLOUDY runs, including the effect of the photoionizing background. We use global ultraviolet background model from Faucher-Giguère et al. (2009) that heats and ionizes the gas in an ionization equilibrium approximation. We apply on-the-fly ionization corrections in denser gas to account for the self-shielding based on the local Jeans-length

¹ Project web site: <http://fire.northwestern.edu/>

Table 1. Simulation details. M_h^0 and M_*^0 are the total mass and stellar mass of the largest halo in the simulation at $z = 0$. $R_{1/2}$ is the radius of the region where half of the stellar mass is enclosed. m_b is the mass of a gas particle in the simulation; m_{dm} is the mass of a DM particle in the simulation. ϵ_b is the minimum gravitational smoothing length of gas; ϵ_{dm} is the Plummer equivalent gravitational smoothing length of DM. The simulation name convention is as follows, ‘mXX’ refers to the halo mass $\sim 10^{XX} M_\odot$. ‘d’ are the corresponding DM-only simulations, e.g. dm09 corresponds to m09. DM-only and hydrodynamical simulations have the same initial conditions, except that the gas particles are absorbed into the DM particles in DM-only simulations. m09, m10, m11, m12v, m12i and m12q are from H14, whereas m10h1297, m10h1146, m10h573 and m11h383 are new simulations first presented in this work.

Name	M_h^0 (M_\odot)	M_*^0 (M_\odot)	R_{vir} (kpc)	$R_{1/2}$ (kpc)	m_b (M_\odot)	ϵ_b (pc)	m_{dm} (M_\odot)	ϵ_{dm} (pc)
m09	2.6e9	4.6e4	36	0.49	2.6e2	2.0	1.3e3	29
m10	7.9e9	2.3e6	50	0.51	2.6e2	2.9	1.3e3	29
m11	1.4e11	2.4e9	1.4e2	6.9	7.1e3	7.1	3.5e4	71
m12v	6.3e11	2.9e10	2.2e2	1.8	3.9e4	21	2.0e5	2.9e2
m12i	1.1e12	6.1e10	2.7e2	4.3	5.7e4	20	2.8e5	1.4e2
m12q	1.2e12	2.1e10	2.8e2	3.6	7.1e3	20	2.8e5	2.1e2
dm09	3.3e9	–	39	–	–	–	1.6e3	29
dm10	9.3e9	–	54	–	–	–	1.6e3	29
dm11	1.6e11	–	1.4e2	–	–	–	4.3e4	71
dm12v	7.7e11	–	2.4e2	–	–	–	2.4e5	2.9e2
dm12i	1.1e12	–	2.9e2	–	–	–	3.4e5	1.4e2
dm12q	1.4e12	–	3.0e2	–	–	–	3.4e5	2.1e2
m10h1297	1.3e10	1.7e7	62	1.8	1.5e3	4.3	7.3e3	43
m10h1146	1.6e10	7.9e7	65	2.5	1.5e3	4.3	7.3e3	43
m10h573	4.0e10	3.2e8	88	3.4	1.5e3	10	7.3e3	1.0e2
m11h383	1.6e11	4.0e9	1.4e2	7.2	1.2e4	10	5.9e4	1.0e2
dm10h1297	1.6e10	–	66	–	–	–	8.8e3	43
dm10h1146	1.8e10	–	68	–	–	–	8.8e3	43
dm10h573	4.2e10	–	91	–	–	–	8.8e3	1e2
dm11h383	1.7e11	–	1.4e2	–	–	–	7.1e4	1.0e2

approximation (to determine the surface density), which provides an excellent match to a full ionization radiative transfer calculation (Faucher-Giguère et al. 2010, 2015; Rahmati et al. 2013).

SF is allowed only in dense, molecular, self-gravitating regions above $n > 10\text{--}100\text{ cm}^{-3}$. This threshold is much higher than that adopted in most ‘zoom-in’ simulations of galaxy formation (the high value allows us to capture highly clustered SF). We follow Krumholz & Gnedin (2011) to calculate the molecular fraction f_{H_2} in dense gas as a function of local column density and metallicity. We allow SF only from the locally self-gravitating molecular gas using the efficiency of 100 per cent per free fall time (the actual SF efficiency is feedback regulated).

Our stellar feedback model includes a comprehensive set of physical mechanisms: radiation pressure, SNe (with appropriate momentum and thermal energy input), stellar winds, photoionization and photoelectric heating as described in H14. We do not tune any feedback model parameters but instead directly use the energy, momentum, mass and metal return based on the output of the STARBURST99 stellar population synthesis model (Leitherer et al. 1999). Our feedback model is implemented within the densest ISM material, yet we do not resort to turning off cooling of SN heated gas at any time.

2.3 Initial conditions and zoom-in method

We adopted a ‘standard’ flat Λ CDM cosmology with $\Omega_0 \approx 0.27$, $\Lambda \approx 0.73$, $\Omega_b \approx 0.045$ and $h \approx 0.7$ for all runs. In order to reach the high resolution necessary to resolve a multiphase ISM and to properly incorporate our feedback model we use the ‘zoom-in’ technique. This places maximum baryonic and DM resolution around the halo of interest in a lower resolution, collisionless box (Porter 1985; Katz & White 1993).

We consider haloes with mass from 10^9 to $10^{12} M_\odot$ at $z = 0$ from the FIRE project (H14). Initial conditions of those haloes are listed in Table 1. The simulations m09 and m10 are constructed using the methods from Oñorbe et al. (2014); they are isolated dwarfs. Simulations m11, m12q and m12i are chosen to match a subset of initial conditions from the AGORA project (Kim et al. 2014) while m12v uses initial conditions from Faucher-Giguère & Kereš (2011, higher resolution versions of the run first presented in Kereš & Hernquist 2009). In addition, we resimulated all of these initial conditions using DM-only N -body simulations with the same Ω_m to have a matched set of simulations with and without baryonic physics for a direct comparison.

To improve halo mass coverage in the regime where cores are prominent, $M_h \sim 10^{10}\text{--}10^{11} M_\odot$, we have also simulated additional four haloes with the same FIRE code, also listed in Table 1. These additional simulations are first time presented in this work. Initial conditions were generated using the MUSIC code (Hahn & Abel 2011). We randomly selected haloes with small Lagrangian regions for resimulation from a $40 h^{-1}$ Mpc box. All particles within $3R_{vir}$ at $z = 0$ are enclosed by the Lagrangian region to reduce contamination (Oñorbe et al. 2014). Haloes m10h1297, m10h1146 and m10h573 have masses between m10 and m11, while m11h383 is slightly more massive. They are all isolated field dwarfs.

2.4 Convergence radius

We adopted the method described in Power et al. (2003) to calculate a conservative limit for the convergence radius of the DM profiles in the N -body-only simulations. They found that effective resolution is related to radius where the two-body relaxation time, t_{relax} , becomes shorter than the age of the Universe t_0 . They verified this with

N -body simulations and found out that for this particular problem, well-resolved regions of haloes require $t_{\text{relax}} > 0.6t_0$. At smaller radii, even if the dynamics are locally well resolved, small N -body effects can, over a long integration time, artificially turn cusps into cores. Given the enclosed number of particles N and the average density of the enclosed region $\bar{\rho}$ one can show that

$$\frac{t_{\text{relax}}(r)}{t_0} = \frac{\sqrt{200}}{8} \frac{N}{\ln N} \left(\frac{\bar{\rho}}{\rho_{\text{crit}}} \right)^{-1/2}, \quad (1)$$

where ρ_{crit} is the critical density. We define r_{Pow} as the smallest radius that fulfils $t_{\text{relax}} > 0.6t_0$ for the DM-only simulations and use r_{Pow} to conservatively estimate where DM profiles are converged. We show the minimum particle mass required for converged profiles at 0.3–3 per cent R_{vir} in Appendix A and discuss implications and limitations of this convergence criterion in Section 5.4.

2.5 Halo finding

We identify haloes and estimate their masses and radii using the Amiga Halo Finder (AHF) (Knollmann & Knebe 2009).² AHF uses an adaptive mesh refinement hierarchy to locate the prospective halo centre (Knebe, Green & Binney 2001). We use the Bryan & Norman (1998) formulae to determine the virial overdensity and virial radius, R_{vir} , and quote the halo mass as the mass enclosed within the virial radius. We follow the main progenitor of a halo using the merger tree code included in AHF and check the growth history of each individual halo, making sure that we follow the same main progenitor. We use the histories of the main progenitors to study the time evolution of the density profile. Occasionally, AHF misidentifies the main progenitor of the halo, so sometimes R_{vir} temporarily decreases. We avoid this problem by searching for another halo with larger R_{vir} and its centre within 50 kpc of the centre of the main progenitor in the previous snapshot.

During ongoing mergers and for galaxies with large clumps of stars and gas, AHF adaptive centring on the highest overall density might quickly change over time and might not centre on the stars. This is especially important in haloes with shallow DM profiles and relatively shallow distributions of stars. To avoid this issue and have a consistent centring on the stellar component, we use two-step procedure to identify the centre. First we use AHF to define R_{vir} and an approximate centre. Within $0.1R_{\text{vir}}$ around this centre we place the stellar mass volume density on a grid. We search for progressively higher overdensities which we enclose in an isodensity ellipsoid. Once this ellipsoid includes less than one quarter of the total galaxy stellar mass we stop the procedure. The centre of the ellipsoid is our new halo centre. This ensures consistent centring of all profiles on the galactic stellar distribution. We tested this procedure and found that our newly defined centre is closer to the DM density peak than the original AHF for cases with shallow central cores and shows less stochastic variations of the central density slope over time.

3 DARK MATTER PROFILES AND THEIR TIME EVOLUTION

The focus of this paper is the effect of stellar feedback on the DM distribution in our simulated haloes. There are three major areas we explore: the relation between the halo mass and the central DM profile, the time evolution of the inner DM density and the changes to galaxy structure caused by the effects of stellar feedback.

3.1 Dark matter profiles

Fig. 1 shows spherically averaged DM density profiles of six simulated haloes at $z = 0$. We focus on the inner regions of haloes, 0.002 – $0.2R_{\text{vir}}$ where galaxies reside and where effects of feedback are expected to be measurable. We show both the profiles from DM-only simulations as well as DM profiles from simulations with baryonic physics. DM-only profiles are renormalized to account for the lower global Ω_{DM} in simulations with baryons. Effects of the baryonic physics are visible in most haloes to a different degree and are the largest in m11 and m12v simulations. In m09 the DM density profile is almost the same in simulations with and without baryons, while in m10, a small, resolved core forms in the central region. The density profile in m11 has the largest core and the lowest central DM density. In m12v, central density starts increasing again and the relative core size decreases, but differences in profiles are present all the way to several percent of R_{vir} . In the two most massive haloes we analyse, m12i and m12q, differences in central region are even smaller, although profiles are still shallower than what is expected based on N -body simulations. However in Section 3.5, we show that the effect of feedback in these haloes is significant and largely cancels out the gravitational influence of baryons that is expected to steepen the profile seen in DM-only case. We note that all of the plotted range is resolved with many gravitational softenings of the DM particles. We also show the more conservative Power convergence criterion which is typically a fraction of a percent of R_{vir} for all of our haloes.

We quantify the effect of feedback on DM distribution using two parameters that are frequently estimated from observations: the central slope α of the DM density profile ($\rho_{\text{DM}} \propto r^\alpha$) and the core radius r_{core} of the pseudo-isothermal sphere (see equation 2). Examples of the fits are shown in Fig. 2.

3.2 Inner slopes of dark matter halo profiles

We estimate the slope α of the DM density profile by fitting a power-law relation $\rho_{\text{DM}} \propto r^\alpha$ in the 1–2 per cent R_{vir} interval. This range is well resolved for all of our main haloes at $z = 0$ and it is physically meaningful as it shows the relative profile change at a fixed fraction of the halo size. For dwarf galaxies, this is close to the region where observations indicate shallow and core-like profiles in low-mass galaxies (typically measured at a few hundred pc; see Walter et al. 2008; Oh et al. 2011, 2015; Hunter et al. 2012). We also show slopes at 0.5–1 per cent R_{vir} for comparison. Example of the fit for α are shown in Fig. 2. We have also measured α in the fitting ranges of 0.3–0.7 and 1–2 kpc. In Appendix B we discuss limitations of these alternative fitting ranges and show that general trends of α with halo mass are similar to our default fitting choice.

Fig. 3 shows α as a function of the halo virial mass, M_h , at different redshifts. We only show main haloes with more than 10^5 DM particles and remove all subhaloes and haloes with more than 1 per cent contamination by more massive DM particles within the inner $0.1R_{\text{vir}}$.³ The open circles represent α whose fitting range contain regions smaller than $0.5 r_{\text{Pow}}$ and/or larger than $1/3 r_s$, where r_s is the scale radius of the NFW profiles.⁴

³ Note that in most runs we use a ‘padding region’ around our zoom-in region where mass resolution is lower only by factor of 8. Mild contamination with such particles can sometimes occur within R_{vir} but typically has no consequence on the evolution of halo gas or central slopes.

⁴ We estimate r_s from the concentration $c = R_{\text{vir}}/r_s$ at a given mass, from the concentration–mass relation of Dutton & Macciò (2014).

² <http://popia.ft.uam.es/AHF/Download.html>

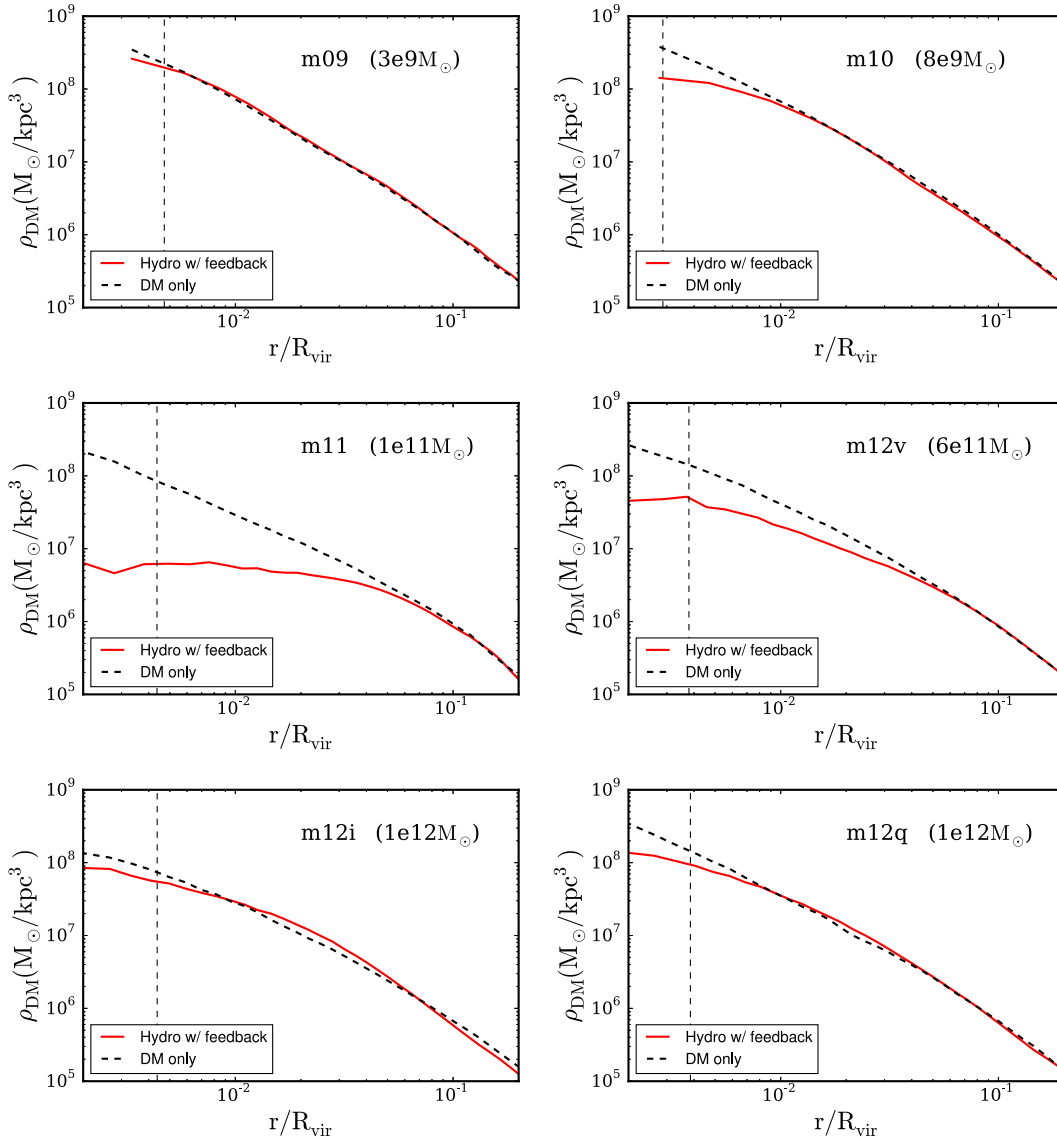


Figure 1. DM density profiles of haloes at $z = 0$. Black dashed lines represent collisionless DM-only simulations; red solid lines represent simulations with baryons and stellar feedback. The Power radius r_{pow} , within which N -body relaxation effects can become important, is shown with vertical black dashed lines. The halo masses are shown in the brackets. Baryonic feedback reduces the central DM density, especially at around $M_h \sim 10^{11} M_\odot$.

We focus on $z \leq 2$ when profiles of haloes start to stabilize as rapid halo growth subsides. At $z = 0$, the simulated haloes show a clear tendency to form shallow central profiles at $M_h \sim 10^{10} - 10^{11} M_\odot$. All of the profiles in this range are significantly shallower than the NFW profile. More accurate estimate of the halo mass and stellar mass ranges where feedback flattens central slopes will require a larger number of simulations as our statistic are currently limited. When profiles are measured at even smaller radii, 0.5–1 per cent of R_{vir} , profiles are typically even more shallow. At $z = 2$ we see that the scaling with mass shows much larger dispersion, which owes to very bursty SF and central halo regions that are just coming out of the fast growth stage. We later show that in intermediate mass haloes at a fixed physical radius, DM profiles get shallower with time.

It is interesting to notice that low-mass dwarfs with $M_h \ll 10^{10} M_\odot$ do not develop density cores even at 1 per cent of R_{vir} (which is typically only several hundreds of parsecs). As we

discuss later, only a small fraction of baryons are converted to stars in these haloes, owing to efficient feedback and effects of the UV background. The energy available from a small number of SNe is not sufficient to dramatically modify the DM distribution. Around $M_h = 10^{10} M_\odot$, the slope of the inner density profile increases rapidly with mass, indicating the development of DM cores. This seems to be a ‘threshold’ halo mass needed to develop large cores. As discussed in Oñorbe et al. (2015), small differences in SF histories in haloes close to this threshold can result in the substantial difference in central slopes of the DM distribution.

Finally, in haloes with mass comparable to the Milky Way (m12v and m12i) profiles steepen again and are only slightly shallower than NFW. These haloes have deep potential wells that can retain a large fraction of available baryons and convert them into stars. Baryons are actually expected to steepen the DM profiles to $\alpha < -1$ owing to adiabatic contraction of DM. However, bursty feedback largely cancels and in some cases even overcomes this expected

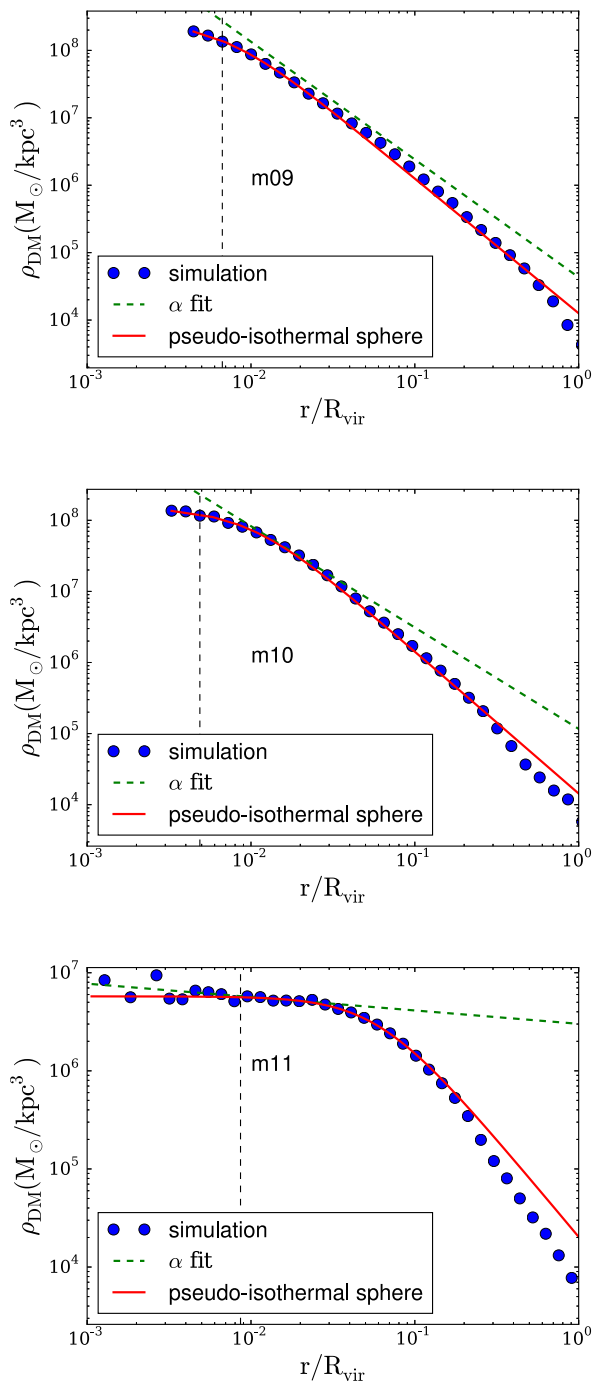


Figure 2. DM density profiles of haloes at $z = 0$ from hydrodynamic simulations (blue filled circles), fits with the pseudo-isothermal sphere (black solid line) and fit with a power-law model ($\propto r^\alpha$) at 1–2 per cent R_{vir} (green dashed line). The black vertical line shows the convergent radius according to the Power criterion. The pseudo-isothermal sphere is a good fit to the central regions of the simulated haloes and provides a good estimate of the core sizes.

effect of contraction, resulting in slopes $\alpha \gtrsim -1$. The interplay between baryonic contraction and stellar feedback will be discussed in Section 3.5.

Fig. 4 shows the scaling of profile slope with galaxy stellar mass (lower panel) and M_*/M_h (upper panel). In terms of stellar mass, feedback significantly modifies DM slopes in the

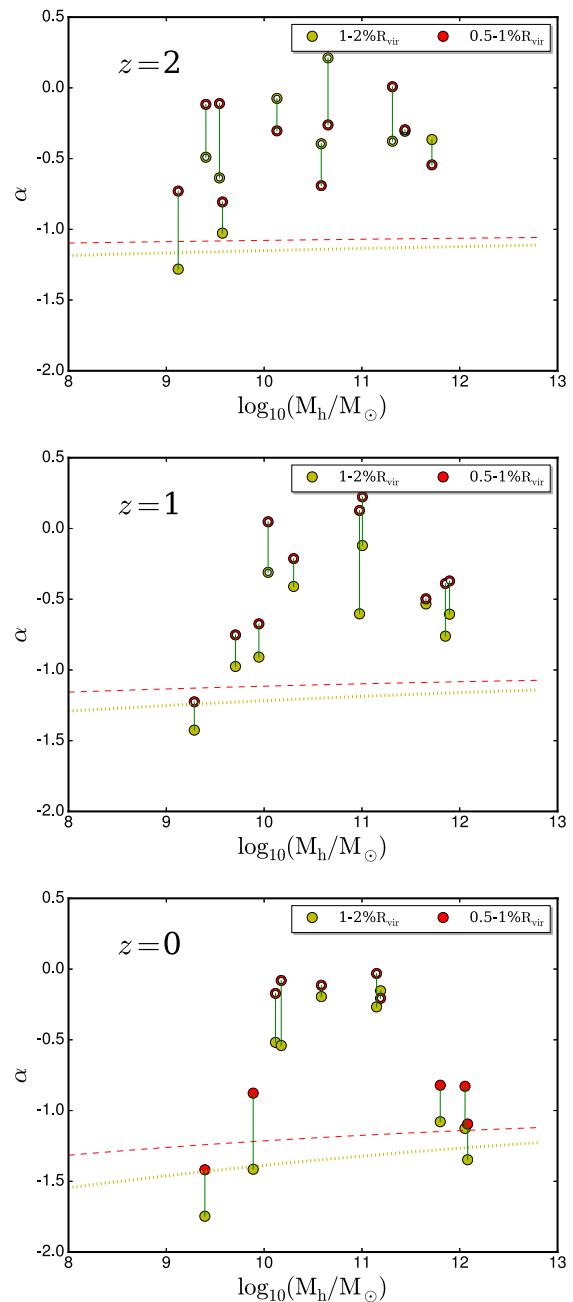


Figure 3. Slopes of DM density α as a function of halo mass at different redshifts. The green dotted and red dashed lines represent the expected α for the NFW profiles measured at 1–2 and 0.5–1 per cent R_{vir} , respectively. The concentration of the NFW profiles is evolved with redshift as in Dutton & Macciò (2014). Filled circles represent α for simulated haloes in which the fitting range is larger than $0.5 r_{\text{pow}}$ and smaller than one-third of r_s . Open circles represent the slopes in haloes in which at least one of these criteria is not satisfied (see the main text for other selection criteria). At $M_h \sim 10^{10} - 10^{11} M_\odot$, baryonic effects lead to profiles significantly shallower than the corresponding NFW profiles from N -body simulations.

$M_* \sim 10^7 - 5 \times 10^9 M_\odot$ range, with a fast transition from cusps to cores occurring at a few times $10^6 - 10^7 M_\odot$. Overall trends of α with M_h in Fig. 3 are similar to the result of Governato et al. (2012) and (Di Cintio et al. 2014, see also a recently submitted work by Tollet et al. 2015). However, we stress that both of these simulations simply suppress cooling in dense gas after SNe explosions rather

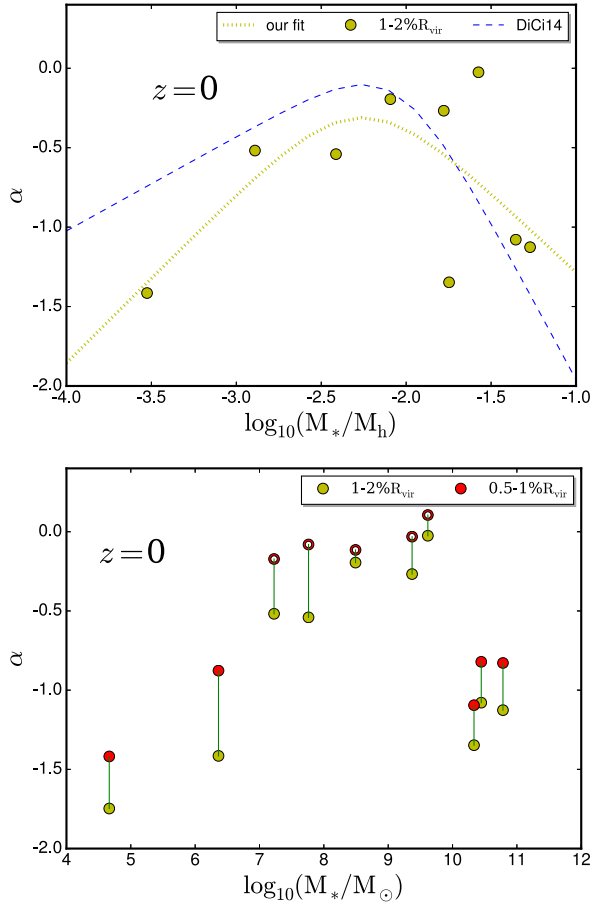


Figure 4. (Upper) Relation between α and ratio between M_* and M_h . The blue dashed line show the fit from Di Cintio et al. (2014) whereas the yellow dotted line shows the same fitting for our data. (Lower) Relation between α and M_* . The symbols are explained in Fig. 3. The DM profiles near halo centres are cuspy at the lowest and highest masses, and shallowest at $M_* \sim 10^8\text{--}10^9 M_\odot$ and $M_*/M_h \sim 0.01$.

than explicitly treat most of the feedback processes around young stars. Furthermore the spatial and mass resolution is typically better in our simulations, by about a factor of 10 in mass. This leads to some differences in the slopes of DM haloes that are illustrated in the upper panel of Fig. 4. In general, profile slopes increase faster with M_*/M_h compared to the previous ‘subgrid’ models, suggesting faster transition from cusps to cores. We caution that a small number of simulated haloes in both samples could also be responsible for some of the differences. In addition we find that the central slope relation is different for the inner 1 per cent or inner 2 per cent of R_{vir} , which means that the fitting formula in Di Cintio et al. (2014) is not generally applicable. We compare our result with observations in Section 4.3.

3.3 Core radii

In addition to the inner slope, we also examine another parameter, the core radius r_{core} of the halo. We quantify the core size using the pseudo-isothermal sphere fit that is frequently used to describe DM density profiles (e.g. Begeman 1987; Broeils 1992; de Blok & McGaugh 1997; Verheijen 1997; another popular fit is the Burkert

Table 2. The core radii of the best-fitted pseudo-isothermal spheres (equation 2) of the simulated haloes at $z = 0$. Large cores of 3–4 per cent R_{vir} form at $M_h \sim 10^{10}\text{--}10^{11} M_\odot$.

$z = 0$	r_{core} (kpc)	$r_{\text{core}}/R_{\text{vir}}$
m09	0.17	0.0048
m10	0.38	0.0073
m11	4.7	0.034
m12v	1.4	0.0061
m12i	2.0	0.0073
m12q	1.2	0.0043
m10h1297	2.0	0.032
m10h1146	2.1	0.033
m10h573	3.6	0.041
m11h383	5.7	0.041

profile, e.g. Salucci & Burkert 2000). Density profile is given by

$$\rho(r) = \rho_0 \left[1 + \left(\frac{r}{r_{\text{core}}} \right)^2 \right]^{-1}, \quad (2)$$

where ρ_0 is the central DM density. We use equation (2) to fit the spherically averaged DM density profiles of the simulated haloes. The two free parameters, r_{core} and ρ_0 , were determined through a χ^2 minimization fitting procedure starting at $r = 0.1$ kpc and ending at $r = \min[R_{\text{vir}}, 100 \text{ kpc}]$. Table 2 lists r_{core} for all haloes analysed in this work. Examples of our fits are shown in Fig. 2. In general fits agree well (to better than few tens of percent within $0.1 R_{\text{vir}}$) with the DM density profiles for all haloes with $M_h < 10^{12} M_\odot$. For m12q and m12i pseudo-isothermal profiles deviate significantly from the DM distribution. Evidence of cores is present in all haloes. The core size is smallest (relative to R_{vir}) in the m09 run (< 0.5 per cent R_{vir}) and largest in m11h383 and m10h573 (> 4 per cent R_{vir}), where we also find the shallowest central slope. Cores of the size of $> 0.005 R_{\text{vir}}$ are present even in Milky Way-mass haloes, albeit with higher central DM density and less shallow central slopes than at $M_h \sim 10^{11} M_\odot$. We compare the core radii from our simulations to observations in Section 4.3.

3.4 Time evolution of α

Next, we investigate the time evolution of the central slope of the DM distribution, for five representative haloes from our sample. Left-hand panels of Fig. 5 show the time evolution of α measured at 1–2 per cent of the halo virial radius at $z = 0$ ($R_{\text{v}0}$). For each halo, this radius is kept fixed in physical units at all times.⁵ In m10, $\alpha < -1$ at all times, which is consistent with its relatively small core size, below 1 per cent R_{vir} . In m10h1297, α is steadily rising from -1 at $z \gtrsim 1$ to ~ -0.5 at $z = 0$. In m11, $\alpha \sim -1$ early on, increasing to $\alpha \sim 0$ around $z \sim 1$ and stays quite flat until late times.⁶ In m12i

⁵ We have also analysed the results within a fixed fraction of the time-dependent virial radius (R_{vir} instead of $R_{\text{v}0}$) but the correlation between stellar feedback and the enclosed DM mass is difficult to interpret because enclosed DM mass increases with R_{vir} .

⁶ At late times $z < 0.5$, this halo undergoes several episodes of very fast central slope variations. We examined this system closely and found that these are caused by close passages of a substructure in an ongoing merger. Vertical lines in Fig. 5 indicate the times of closest passages: they correlate well with temporary drop and strong oscillations in central slope. The close passages can affect the accuracy of locating centre of the galaxy in AHF, which further motivates our two-step centre finding procedure described in Section 2.5.

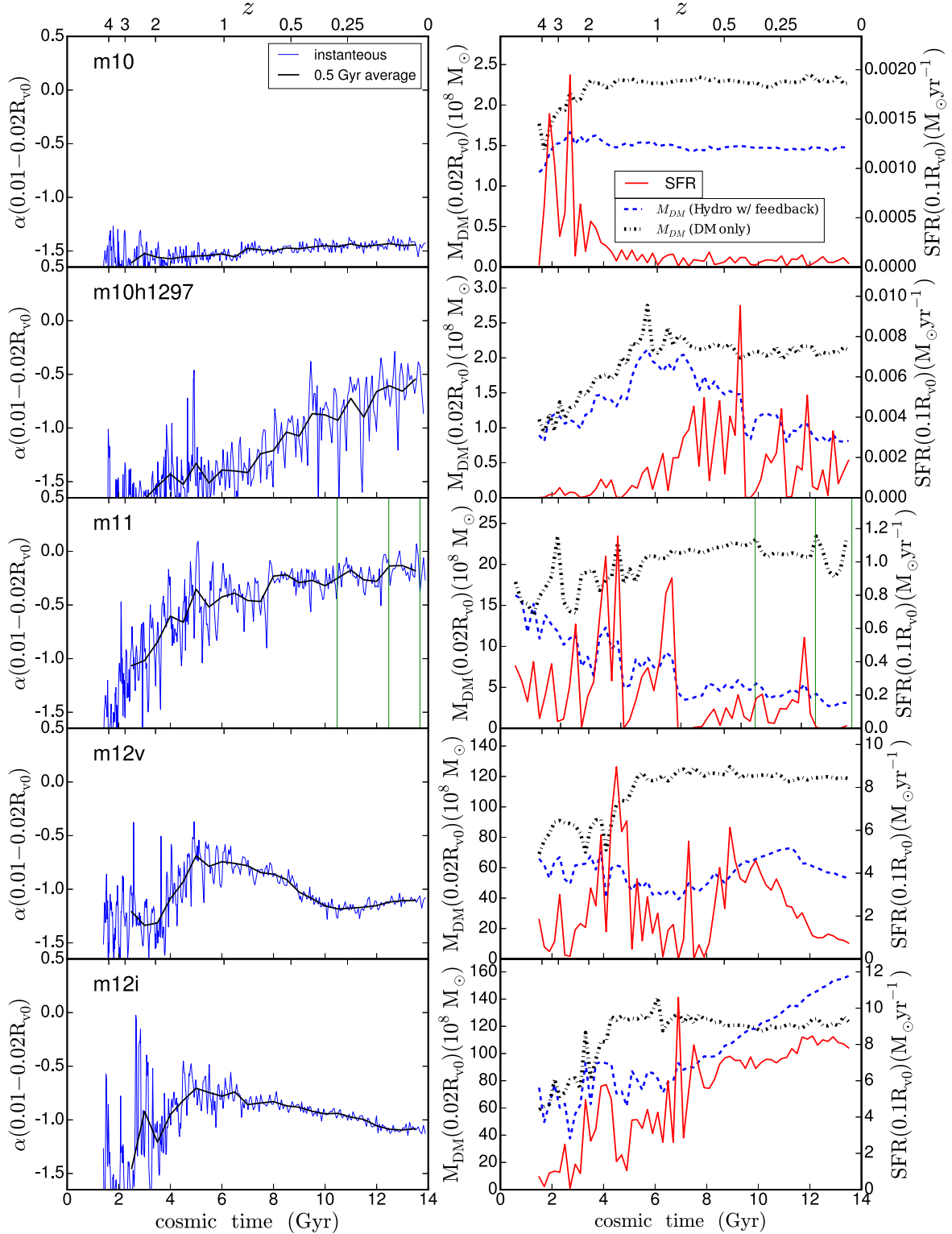


Figure 5. (Left) DM density slope, α , within 1–2 percent $R_{vir}(z=0)$, R_{v0} , as a function of redshift/cosmic time in m10, m10h1297, m11, m12v and m12i, respectively (from top to bottom). The slope is measured at fixed physical radius at all times. Blue lines shows the variations on the time-scale of our simulation output (typically 20–30 Myr) while the black lines show the average values over 0.5 Gyr. (Right) Time evolution of the enclosed DM mass within $0.02R_{v0}$ for DM-only (black dash-dotted) and hydrodynamical simulations with feedback (blue, dashed) and SFR within $0.1R_{v0}$ (red solid), all averaged over 0.2 Gyr. Green vertical lines in m11 panels show the times of close passages of a subhalo. Cores form when the central DM accretion stops but the SF is still bursty, as seen in m10h1297 and m11.

and m12v central DM profiles are flattened to $\alpha \sim -0.7$ at $z \simeq 1$ but steepen afterwards such that central DM slope is $\alpha \sim -1$ at the present time.

The right-hand panels show the time evolution of the enclosed mass within $0.02R_{v0}$ and the star formation rate (SFR) within $0.1R_{v0}$ averaged over 0.2 Gyr.^{7,8} A dense central concentration of DM builds up early in m10, with some fluctuations during the bursty SF epoch at $2 < z < 4$, but as SF subsides the amount of DM in the central region remains almost unchanged until the present time. The correlation between SF and strong outflows of gas that follow the burst (Muratov et al. 2015) and the decrease of DM in the central region is clearly visible in the m10h1297 and m11 panels. Removal of DM mass occurs after strong bursts of SF. We examine this more closely in Section 5.2. While some of the DM gets re-accreted, the central concentration of DM remains lower after the burst for at least several Gyr.

Each strong burst of SF reduces central density of DM, so the final slope and core size are a consequence of several burst episodes over a Hubble time. The overall effect is small in m10, because the SFR decreases to very low values at $z < 4$. From the comparison of DM-only and feedback simulations it is clear that a small difference in central DM concentration was established early on and stays largely unchanged until late times.

A small, fluctuating and early decrease in the DM concentration is also seen in more massive haloes, e.g. m10h1297 and m11. However, the amount of DM in a central region only decreases significantly once the central region finishes rapid growth, at $z \sim 1$ in m10h1297 and $z \sim 3$ in m11. After this stage, DM-only simulations show an approximately constant amount of DM in the central region while baryonic simulations successfully evacuate a large amount of DM from the centre. Unlike m10, they have several ongoing bursts of SF after the rapid-build-up stage, each of which removes a significant amount of DM. It appears that having strong bursts of SF and outflows *after* the inner halo build-up slows down is the key to produce a long-lasting shallow DM density profile. At early times, during the fast build-up stage, shallow profiles are not fully established as showed in Fig. 5.

In m12i and m12v, removal of DM after the peaks of SF is also seen at $z > 1$ when fluctuations are large. After $z \sim 1$, the SF continues at modest level without rapid bursts. At the same time, the enclosed DM mass grows slowly. In these massive haloes re-accretion of DM in the centre occurs when the SFR is low and hierarchical assembly is slow. To explain this we followed the central accumulation of baryonic material and found out that in both haloes baryons start to dominate central mass at $z \sim 1.5$. As a consequence DM gets contracted, increasing the amount in the inner halo (see Section 4). This effect is stronger in the m12i simulation that is more baryon dominated, and accumulates baryons faster at late times.⁹

The core forms after multiple starbursts rather than one single blow-out, which is consistent with the mechanism discussed in

Pontzen & Governato (2012, see also Ogiya & Mori 2014). After a blow-out some re-accretion of DM does occur but even after several Gyr the amount of DM enclosed in central region does not return to the pre-burst level, indicating that the effect on DM is long lived. These trends are the most obvious at $M_h \sim 10^{11} M_\odot$, i.e. in m11. However, in more massive systems we see that cusper profiles are rebuilt at later times. While SF proceeds to late times, it is often spread throughout the disc. In a companion paper (Muratov et al. 2015), we show that at late times the SF activity is not able to eject large quantities of material from galaxies, which is why this continuous SF does not ‘heat’ and remove DM from the centre. Even if the assembly of central regions of the DM haloes slows down at relatively early times, increase in central concentration of baryons at late times can rebuild DM cusps via adiabatic contraction. Our simulations suggest that the contribution of minor mergers to the regrowth of cusps (e.g. Dekel, Devor & Hatzroni 2003; Laporte & Peñarrubia 2015) at late times is likely subdominant or negligible, but a larger number of haloes is needed to confirm this for haloes with diverse growth histories.

We have already seen that core formation depends strongly on halo mass but also on the presence of significant SF episodes after the central region stops accreting DM. This is critical at halo masses where large cores start to develop. We showed this directly by using the initial condition of m10, simulated with a slightly different local feedback coupling scheme that increases late time SF and its burstiness: the outcome is a larger DM density core and a shallower centre slope compared to the original FIRE m10 simulation analysed here (see Oñorbe et al. 2015). This means that if there are bursts of SF activity occurring in dwarf galaxies/haloes around this mass at late times they could result in shallow density profiles by present time.

The SF history of m10 is significantly different from m10h1297. It forms most of the stars before $z = 2$ and becomes passive at late time whereas other haloes remain actively star forming at present. From the observations of dwarf galaxies in the Local Group, Weisz et al. (2014) show that a large fraction of dwarf galaxies have active late time SFs, especially for field galaxies and galaxies with $M_* > 10^6 M_\odot$. Therefore, the SF histories of our simulated galaxies are well within the observed range. Our limited statistics suggest that at $M_* \sim 10^6 - 10^7 M_\odot$ SFR history is closely linked to formations of large cores (see also Oñorbe et al. 2015).

3.5 Halo expansion or baryonic contraction?

In the previous section, we investigated the shapes of profiles of our simulated haloes. Here we examine the effect of a central concentration of baryons on the DM profiles.

First, we examine the net effect of halo expansion via feedback and halo contraction owing to central concentration of baryons. Feedback is dominant in shaping flatter profiles in lower mass haloes, but baryonic contraction largely cancels the feedback effect in Milky Way-mass haloes, such that their final profiles are only slightly shallower than the NFW profile.

In order to estimate the effect of baryonic contraction on the DM profiles, we follow Blumenthal et al. (1986) to calculate the final DM mass distribution M_x , given the final baryon mass distribution m_b and the initial total mass distribution M_i :

$$r[m_b(r) + M_x(r)] = r_i M_i(r_i) = r_i M_x(r)/(1 - F_b), \quad (3)$$

where F_b is the fraction of dissipational baryons, r_i and r are initial and final orbital radii, respectively. We follow the simple semi-analytic model (Dutton et al. 2007), and assume that initial total

⁷ Newly formed stellar particles can move quickly between simulation outputs during the SF burst, owing to feedback induced mass redistribution, so we integrate SF within this larger radius.

⁸ Results averaged over 0.1 Gyr or shorter time-scales are qualitatively similar but show more rapid fluctuations in subcomponents of longer bursts so we selected 0.2 Gyr for the sake of clarity.

⁹ In m12v the SFR is low and the amount of baryons in the central region changes very slowly in the final several Gyr. The enclosed DM at those times is also affected by ongoing minor mergers that cross very close to the centre and are later disrupted causing variations in the enclosed density.

halo density profile is the one from our DM-only simulation, and use the final distribution of baryons (stars and gas) in the full physics simulation to estimate the contraction. Therefore, we use m_b from hydrodynamical simulations with stellar feedback, and M_i is the halo mass from DM-only simulations.¹⁰

Fig. 6 shows density profiles from collisionless simulations, full feedback simulations and from models with a contracted DM halo. Left-hand panels show three different DM profiles for each halo and baryon density profiles from feedback simulations for reference. Right-hand panels show the total density profiles including DM and baryons (gas and stars).

In m10, the estimated effect of adiabatic contraction is very small because the fraction of baryons in the centre of m10 is small. The baryon density is less than one-tenth of DM density near the centre. However, the feedback slightly expands the DM and forms a small core.

In m10h1297 and m11, stellar feedback strongly affects DM distribution, making it much shallower than in the corresponding collisionless run. It is interesting that in m10h1297, feedback significantly flattens the DM profile, creating a large core, but the halo remains DM dominated at all radii. In m11, the baryon density is a slightly larger fraction of the total, but still significantly lower than the corresponding DM density in the collisionless simulation. The contracted profiles for both of these haloes are therefore very similar to the DM profiles from the corresponding collisionless simulations.

In more massive haloes, the m12 series, baryonic contraction is expected to significantly steepen the DM profiles, because their central regions are baryon dominated. While m12v shows a strong effect of feedback, the profile of m12i is relatively steep in the centre, similar to the NFW profile. However, when compared to the expectations of baryonic contraction we see that the resulting profile is much shallower than contracted NFW halo for all of the plotted simulations. This demonstrates that even in these massive haloes, feedback has a strong effect on shaping the final DM profiles, and largely cancels out the effect of contraction. In general the expansion of the halo by the stellar feedback causes an order of magnitude difference in the density profiles around 1 kpc in $10^{11-12} M_\odot$ haloes when compared to expectations of a simple baryonic contraction model.

While feedback effects on the DM distribution are substantial even in m12 series haloes, in those haloes baryons completely dominate the central few kpc at $z = 0$. This is why the differences between the total matter density profiles in simulations with feedback and profiles expected from the contracted original DM halo are relatively small (right-hand panels). We will return to these total matter profiles shortly and show that even this small effect has measurable consequences in the circular velocity curves of galaxies. It is interesting to note that the total matter distribution in the inner 20 per cent of the R_{vir} of m12 series is well approximated by the isothermal density profile ($\rho \propto 1/r^2$).

The results from our m12v simulation are consistent with the strong feedback run in Macciò et al. (2012b) who showed core formation in a $7 \times 10^{11} M_\odot$ halo. We do however find slightly higher central density of DM than reported in Macciò et al. (2012b). While this might be just a matter of small number statistics, our simulation results should give more accurate predictions for the central profiles

both because of more realistically implemented stellar feedback and because of the higher resolution. At masses similar to the Milky Way ($M_h \sim 10^{12} M_\odot$), the DM density distribution in the centre is only slightly shallower than the NFW profile as strong feedback effects and adiabatic contraction of such expanded halo nearly cancel out.

4 OBSERVABLE CONSEQUENCES

The following subsections show how our simulations with stellar feedback can alleviate the tensions between previous simulations and observations, including the ‘lack’ of baryonic contraction, the ‘TBTF’ and *cusp/core* problems.

4.1 Rotation curves and the Tully–Fisher normalization

The distribution of matter in galaxies can be measured with the rotation curves. For disc galaxies, there is a tight relation between their luminosity (or mass) and their circular velocity, the so called TFR (Tully & Fisher 1976). Here we examine the effect of stellar feedback and baryonic contraction on rotation curves.

Fig. 7 shows the rotation curves of haloes from simulations and baryonic contraction calculations for m11, m12v and m12i. Here $V_c = \sqrt{GM(r)/r}$. The rotation curves of simulated galaxies do not show profiles expected from the NFW haloes affected by adiabatic contraction. Instead, galaxies have lower mass concentrations in the centres, resulting in lower circular velocities. Therefore feedback effectively prevents build-up of high densities expected from strong adiabatic contraction. As expected, the rotation curves in the m12 series are approximately flat at radii larger than several kpc, while for the massive dwarf m11 the rotation curve is rising.

The effect of feedback, which cancels the effect of baryonic contraction, turns out to be very important for the normalization of the TFR. We do not perform a detailed comparison with observations, which would require mimicking observational measurements of the rotational velocity and luminosity. We plan to study this in future work. Here, instead, we focus on the relative effect of the feedback on DM profiles that determine the normalization of TFR. To show this effect and compare our simulated galaxies with the observed TFR, we measure the circular velocity of the halo $V_c = \sqrt{GM(r)/r}$. We measure circular velocity at 2.2 ‘disc scale length’ $V_{2.2}$, approximately mimicking frequent observational approach (Dutton et al. 2010). We first measure the half-mass radius of the stellar distribution $r_{1/2}$ and use the relation between half-mass radius and scale radius of an exponential disc to define disc scalelength as $r_d = r_{1/2}/1.67$.

Fig. 8 shows the TFR of main galaxies with $10^9 < M_*/M_\odot < 10^{11}$, and the best fit of the observed TFR from Dutton et al. (2010). They derived the TFR using the data from Courteau et al. (2007) and with the best fit

$$\log_{10} \frac{V_{2.2}}{[\text{km s}^{-1}]} = 2.064 + 0.259 \left(\log_{10} \frac{M_*}{[10^{10} M_\odot]} \right). \quad (4)$$

This relation was derived for relatively massive galaxies, most with $\log V_c [\text{km s}^{-1}] > 1.8$. Similar relations were found for dwarfs, however, with significantly enlarged scatter and non-uniform way of measuring V_c (e.g. Ferrero et al. 2012). We therefore limit discussion here to galaxies with $M_* > 10^9 M_\odot$.

It is clear that the strong feedback which reduces the effect of adiabatic contraction is a necessary ingredient in reproducing and explaining the TFR in massive galaxies with $M_* > 10^{10} M_\odot$. While there are direct effects of feedback on the distribution of baryons within galaxies, feedback effect on the distribution of DM

¹⁰ We also tried to account for the loss of baryons due to feedback but the mass-loss is much smaller than the total mass, especially in m12 series. The difference in circular velocity with or without the missing baryons is less than a few per cent in m11 and negligible at higher masses.

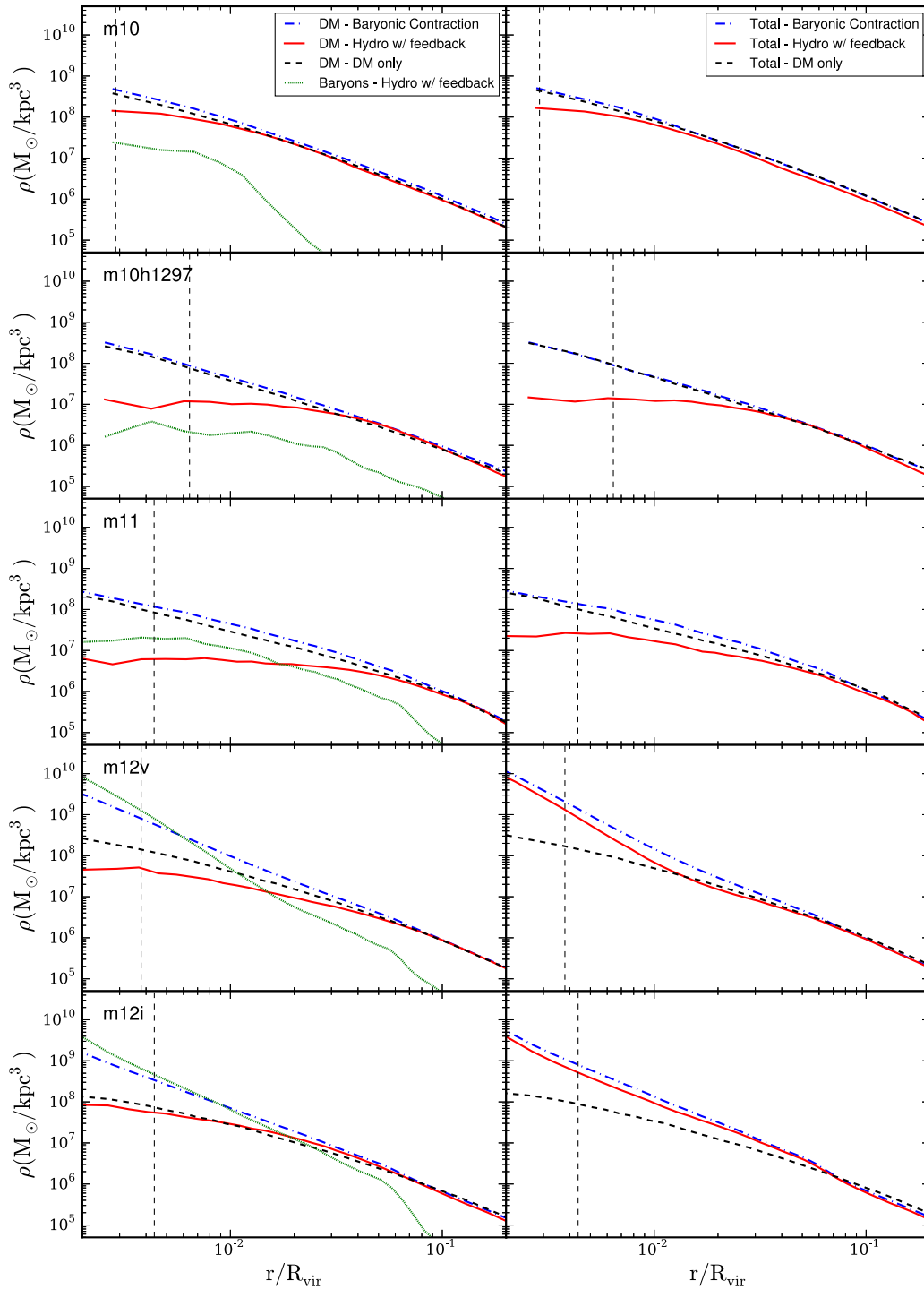


Figure 6. (Left) DM density profiles of m10, m10h1297, m11, m12v and m12i at $z = 0$. Different line colours show the expected DM profile in simulations with baryons and feedback (solid; red), in collisionless simulations (dashed; black) and in calculations including baryonic contraction (dot-dashed; blue). The green dotted line shows the total baryon density, including both gas and stars, in feedback runs. (Right) Total density profiles of the same haloes, including both DM and baryons. The Power convergence radii are shown as dashed vertical lines. In haloes where baryons dominate in the central regions, total and DM densities based on the simple baryonic contraction model are higher than the actual densities in our simulations with baryonic feedback. Feedback effectively cancels the effect of baryonic contraction.

is also an important ingredient in establishing the TFR. Simulated galaxies appear to better match the observed TFR than our model with baryonic contraction. The circular velocities in baryonic contraction calculations are higher by a factor of 1.2–1.5 than V_c in simulations.

Our findings confirm previous conclusions that the lack of effective contraction is necessary to explain the Tully–Fisher relation (Dutton et al. 2007; Macciò et al. 2012b). This also explains why previous generations of simulations without efficient feedback had trouble matching the normalization of TFR (e.g. Steinmetz &

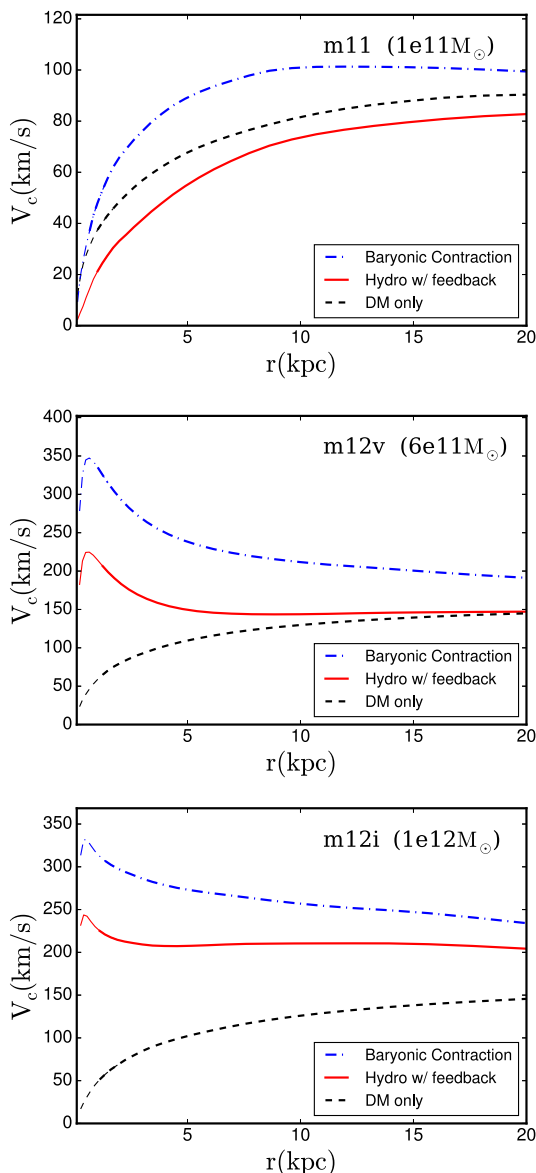


Figure 7. Rotation curves of haloes, including m11, m12v and m12i. Red solid lines represent rotation curves from simulation with feedback, black dashed lines show results from DM-only simulations, whereas blue dot-dashed lines represent rotation curves with baryonic contraction. Dashed lines show the region within the Power radius. Their halo masses are shown in the brackets. The rotation curves in simulations with baryonic feedback are lower than in a simple baryonic contraction model. The simulated Milky Way-mass haloes show flat rotation curves, while the large dwarf galaxy m11 shows a rising rotation curve.

Navarro 2000). In these models too much gas collapsed to the centre exerting strong contraction of DM halo without previously affecting the DM distribution.

4.2 Implication for the ‘Too Big To Fail’ problem

In addition to the cusp/core problem, CDM simulations are also challenged by another problem, the so-called TBTF problem (Boylan-Kolchin et al. 2011; Garrison-Kimmel et al. 2013). The simplest version of this problem is that the observed Milky Way satellite galaxies have much lower central circular velocities than subhaloes from cosmological collisionless CDM simulations. This

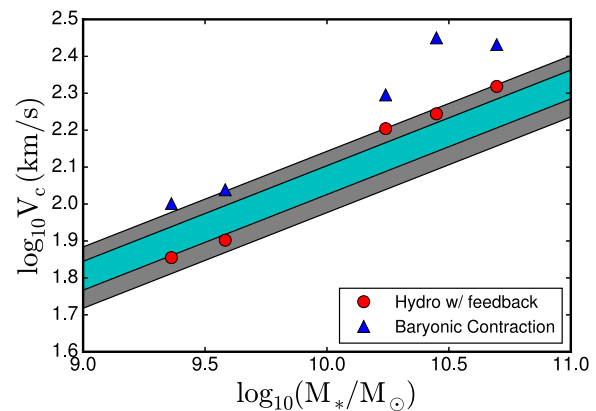


Figure 8. Stellar mass–Tully–Fisher relation for observed galaxies (equation 4) and for the simulated haloes (m11h383, m11, m12v, m12q and m12i). The y-axis is the rotation velocity measured at 2.2 disc scalelength. The shaded regions show equation (4) with 1σ uncertainty ($\sigma = 0.039$) (cyan) and 2σ uncertainty on the zero-point (grey), respectively. Stellar feedback, which counteracts the effects of adiabatic contraction, appears necessary to establish the observed normalization of the Tully–Fisher relation.

either means that massive subhaloes do not have corresponding match in observed satellite galaxies or that central regions of predicted CDM subhaloes are too dense compared to observed haloes. This seems to be generic problem, independent of the halo formation history as similar effects are also observed in the Local Group and for dwarf galaxies in general (Ferrero et al. 2012; Garrison-Kimmel et al. 2014; Papastergis et al. 2015) including non-satellite galaxies.

We have already shown that stellar feedback can reduce the central DM density, which could potentially resolve this discrepancy, without invoking a different type of DM. In Fig. 9, we show the circular velocity profiles of the central kpc of m10, m10h1297 and m10h1146 at $z = 0$ along with their corresponding DM-only simulations. These are our best-resolved systems with galaxy stellar mass $\sim 2 \times 10^6 - 8 \times 10^7 M_\odot$ at $z = 0$, which are close to the stellar masses of the galaxies for which the ‘TBTF’ problem was demonstrated (Boylan-Kolchin et al. 2011; Garrison-Kimmel et al. 2013). For comparison we also include the observational data from Milky Way satellite galaxies (Strigari, Kaplinghat & Bullock 2007; Walker et al. 2009; Wolf et al. 2010) and Local Group field galaxies (Kirby et al. 2014). It is clear that feedback strongly reduces circular velocities in the central few hundred pc with respect to collisionless cold matter simulations, in some cases by a large factor. Such reduced circular velocity implies that observed dwarf galaxies (including satellite galaxies) should not be associated with haloes and subhaloes from DM-only simulations with the same circular velocity; instead these should be connected to the predicted higher circular velocity analogues, whose circular velocity is now reduced owing to feedback. Our results strongly suggest that this effect would dramatically reduce the number of ‘massive failures’ and can alleviate or potentially solve the ‘TBTF’ problem. Our findings qualitatively agree with hints in previous work (Brook & Di Cintio 2015).

It is interesting that the high stellar mass dwarf galaxies (e.g. m10h1146 and m10h1297) have more significant reductions in central rotation velocities (and thus dynamical masses) compared to low stellar mass dwarf galaxies (e.g. m10). This causes the rank order of V_c at small radius (e.g. 500 pc) not to correspond to the rank order of their V_{\max} or their stellar mass, as illustrated in the middle and lower panels of Fig. 9. Direct comparison between DM-only

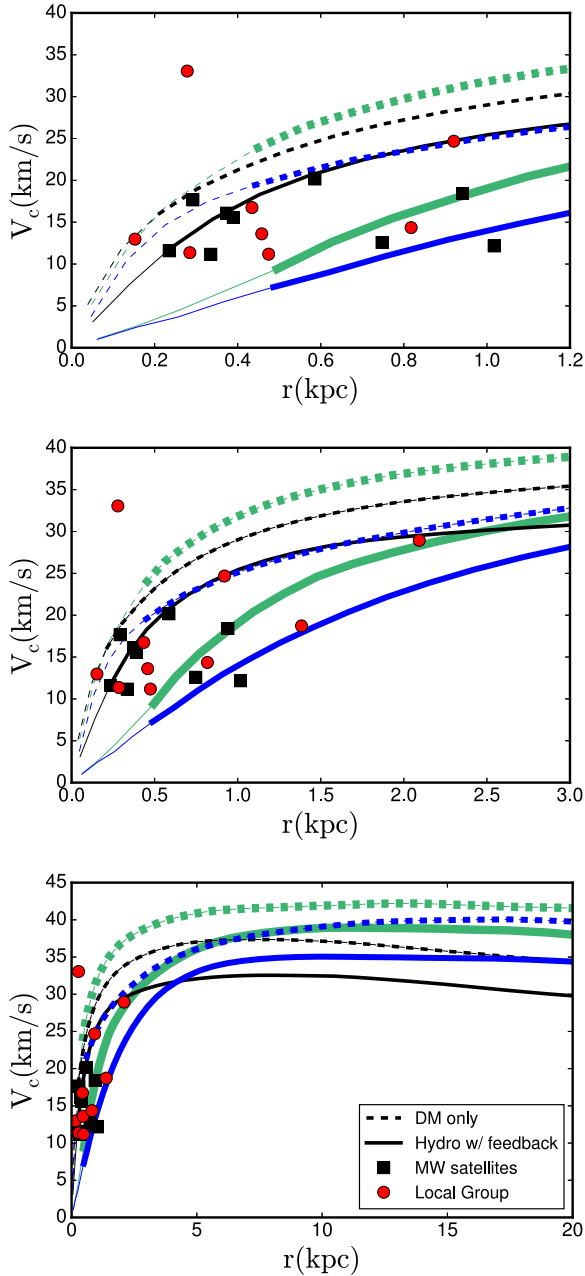


Figure 9. Rotation curves illustrating the BTBF problem, plotted over a range of radial scales. We have included haloes m10 ($M_* = 2.3 \times 10^6 M_\odot$; black, thin), m10h1297 ($M_* = 1.7 \times 10^7 M_\odot$; blue, normal) and m10h1146 ($M_* = 7.9 \times 10^7 M_\odot$; green, thick) and their corresponding DM-only simulations. Thick dashed lines represent the haloes from the collisionless simulations while thick solid lines represent the haloes from the hydrodynamical simulations with feedback. The thin lines show the velocity curves at radius smaller than the Power convergence radius. Black squares show the data from Milky Way bright satellite galaxies (Strigari et al. 2007; Walker et al. 2009; Wolf et al. 2010), while red circles show the isolated dwarf galaxies in the Local Group (Kirby et al. 2014). The panels show three different scales of the same plot to illustrate that the order of rotation curves at small scales may not imply the order of halo masses. Stellar feedback reduces the central circular velocity such that the rotation curves can match the observed dwarfs, suggesting that baryonic feedback may solve the ‘BTBF’ problem. Observational errors are typically smaller than a few km s^{-1} (not shown for clarity).

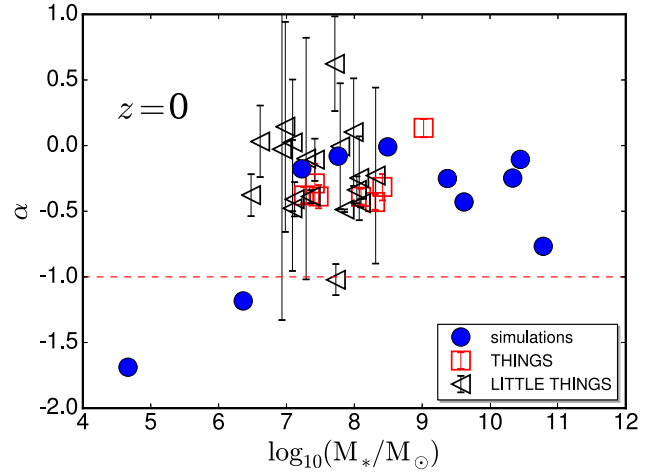


Figure 10. Slope of dark matter density profile, α , from our simulations (measured at $r = 0.3\text{--}0.7$ kpc) compared with α from observations (typically measured at a few hundred pc). Blue filled circles represent the simulated haloes at $z = 0$. Red open squares represent the observed dwarf galaxies from THINGS (Walter et al. 2008; Oh et al. 2011), whereas black open triangles represent the dwarf galaxies from LITTLE THINGS (Hunter et al. 2012; Oh et al. 2015). In the overlapping mass range, the simulated dwarf galaxies have central DM profile slopes in good agreement with the observed dwarfs.

simulations and simulations with baryons is even more complex, and rank order matching of V_c , or V_{peak} from measurements at small radii might lead to incorrect physical interpretations.

Our results only indirectly address the ‘BTBF’ problem in satellite galaxies, because the galaxies we consider here are not satellites but field galaxies. Satellite galaxies of relevant mass in our m12 simulations do not have sufficient mass resolution to study their DM distributions at the galaxy centres. The effect of host galaxies on satellites, e.g. tidal stripping, could also modify the structure of DM haloes (Zolotov et al. 2012; Brooks & Zolotov 2014). However, the effect of lowering the circular velocity of galaxies is generic in the range of stellar masses $2 \times 10^6\text{--}3 \times 10^9 M_\odot$ and we therefore believe that satellite galaxies in this mass range will be affected in the same systematic way.

4.3 Central slopes and core sizes

Oh et al. (2011) and Oh et al. (2015) measured detailed density profiles of field dwarf galaxies, and found flat DM profiles near the centre, in contrast with the cuspy NFW profiles, predicted from N -body simulations. In massive galaxies, such as the Milky Way, baryons dominate in the central few kpc (e.g. Courteau & Dutton 2015) making measurements of central DM properties extremely difficult. We therefore focus on lower mass galaxies/haloes with reliably measured central DM profiles.

In Fig. 10, we show that FIRE haloes at $z = 0$ are in good agreement with the observed slopes of central DM density profiles (see also Governato et al. 2012). This suggests that the inclusion of stellar feedback in our simulations helps resolve the so-called cusp/core problem observed in low-mass galaxies. However, the observed scatter is large and number of simulated objects is limited in the observed mass range. It is therefore clear that much larger sample of model galaxies/haloes as well as more detailed accounting for methodology and selection used in observations are needed to test our model in detail.

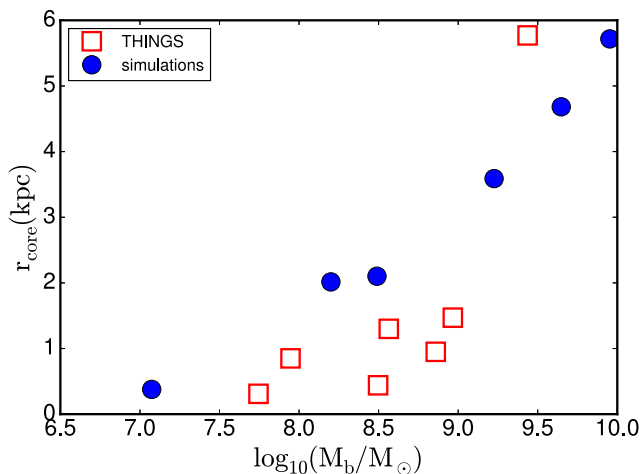


Figure 11. Comparison between the core radii of THINGS dwarf galaxies and our simulations as a function of baryonic mass m_b (see footnote 11) at $z = 0$. Red open squares show the observed core radii (Walter et al. 2008; Oh et al. 2011) while blue circles show the core radii from our simulations. Within the plot range, the core size increases with baryonic mass, which is largely consistent with the observed cores. A larger number of observed and simulated dwarfs is necessary to draw stronger conclusions.

Fig. 11 shows the DM core radii of our simulated galaxies as a function of their baryonic masses and compares them to observations from Walter et al. (2008). In general, for $m_b \sim 10^7\text{--}10^{10} M_\odot$,¹¹ core radius increases with baryonic mass. There is a broad agreement with observed core sizes. A more detailed comparison will require a larger number of haloes and more detailed modelling of the methodology used to calculate the core sizes in observed galaxies.

5 DISCUSSION

We use the FIRE suite of high-resolution cosmological ‘zoom-in’ simulations with physical feedback models to study the properties of the central regions of DM haloes. Our simulations have higher resolution (in most cases the highest resolution to date at a given halo mass), and more explicit and comprehensive implementation of stellar feedback than simulations used previously to study baryonic effects on DM profiles. Critically our models contain no ‘freely adjusted’ parameters tuned to any particular results. We characterize the evolution of halo properties and correlation with galaxy and halo mass and explore the effects of SF driven feedback and adiabatic contraction on the slopes, cores size and circular velocity profiles of galactic haloes.

We find that at $z = 0$ the central slopes of DM density profiles measured at ~ 1 per cent R_{vir} are shallow in haloes $M_h \sim 10^{10}\text{--}10^{11} M_\odot$, but the slopes are cuspy at lower and higher masses. We see a sharp transition around $M_h \sim 10^{10} M_\odot$ from cuspy profiles at lower masses to shallow, core-like profiles at higher masses. Efficient feedback continues to at least $M_h \sim 10^{12} M_\odot$ where the core-like profiles that form during earlier evolution are contracted by baryons at late times into steeper profiles. Final profiles are similar or flatter than NFW, and therefore significantly flatter than expectations for a contracted NFW halo.

Our results are in broad agreement with others found in the literature. For example, we find that feedback is efficient in forming large cores at halo masses $\sim 10^{10} M_\odot$ to few times $10^{11} M_\odot$ which is similar to previous findings (Governato et al. 2012; Di Cintio et al. 2014; Pontzen & Governato 2014). It appears that simulations with bursty SF and outflows, but different small-scale feedback implementations, affect DM profiles in a qualitatively similar way. However, we also found some interesting differences. This is not surprising, given that our simulations are in most cases higher resolution and have explicit treatment of feedback on small scales. For example, we find that in the haloes with large cores, cores are established at around $z \sim 1$ and can grow even at low redshift while other authors find that cores formed very early already at $z = 2\text{--}3$ (e.g. Madau et al. 2014). This could be a consequence of different SF histories that can change core formation (Oñorbe et al. 2015). For example, these authors note that earlier simulations or dwarf galaxies with ‘subgrid’ feedback models produced too many stars at a given halo mass. Alternatively, these differences could indicate that different treatments of material affected by feedback (cooling prevention versus explicit model) cause differences in DM profiles.

We also find cuspy DM profiles at around $6\text{--}7 \times 10^{11} M_\odot$ than the one reported in Macciò et al. (2012b) and therefore weaker halo expansion and somewhat steeper dependence of slope on M_*/M_h than the relation presented in Di Cintio et al. (2014). Overall, better statistics from a larger number of simulated haloes are needed for a more robust analysis of these differences.

5.1 Energetics

Given that the efficiency of conversion of halo gas into stars increases from dwarf galaxies to massive galaxies, it is natural to connect feedback effects to the energy available from stellar feedback. We therefore compare the amount of energy available from feedback with the energy needed to overcome the gravitational potential and move DM outside of the central region. The simplest estimate is to calculate how much energy gets injected into the ISM from the SNe only, which represents a lower limit to the available energy budget.

Previously, Gnedin & Zhao (2002), Ogiya & Mori (2011) and Garrison-Kimmel et al. (2013) claimed that the total feedback energy released in SNe is insufficient to remove enough DM to form large cores. In particular, Garrison-Kimmel et al. (2013) tested if stellar feedback can lower central densities to a degree needed to explain the ‘TBTF’ problem. They simulated SN feedback with time-varying potential and found that the number of SNe needed to match observed profile of a halo hosting $M_* \sim 10^6 M_\odot$ galaxy exceeds the number of SNe produced in most of the dwarf galaxies for the typical initial mass function. However, they did not consider the full growth history of the halo. The frequent mergers and SF bursts at higher redshift could dynamically heat up more DM in the centre, when halo was smaller. Furthermore, they assume mass loading of SNe-driven winds smaller than we find for FIRE galaxies of comparable masses (Muratov et al. 2015). Finally their selected halo has a high concentration compared to a typical subhalo that is expected to host observed dwarf galaxies. As we show below, there is a sufficient amount of energy available to couple to DM in the relevant halo and galaxy mass regime.

Table 3 shows the energy needed to create a constant density core with radius 1 or 3 per cent R_{vir} in haloes from our DM-only simulations and the total SN energy inferred from our hydrodynamic simulations by $z = 0$. We constructed a constant density core in our ‘cuspy’ DM-only simulations by keeping the density profile outside

¹¹ m_b is defined as the total baryonic mass within 20 per cent R_{vir} of the halo.

Table 3. Comparison of the difference in gravitational potential energy ΔU of haloes from DM-only simulation and the same haloes with constant density core of 1 per cent of R_{vir} (left) or 3 per cent of R_{vir} (middle), and their total SN energy (E_{sn} , right), obtained from the corresponding simulations with feedback. Energy input from SNe alone is sufficient to produce cores in m10 and m11.

	ΔU (1 per cent core) (erg)	ΔU (3 per cent core) (erg)	E_{sn} (erg)
m09	5.85e52	6.45e53	2.68e53
m10	1.08e53	2.02e54	2.13e55
m11	3.43e55	6.46e56	2.22e58

the core unchanged and moving the excess mass within the core radius to infinity. Then we calculate the total potential energy for the initial and cored profiles with the formula below:

$$U = -4\pi G \int_0^{R_{\text{vir}}} M(r) \rho(r) r dr, \quad (5)$$

and report the difference, ΔU , in Table 3.

To estimate the SN energy, we assume the energy from one SN $E_{\text{sn}} = 10^{51}$ erg, the fraction of massive stars which can produce SNe $\xi(m_* > 8 M_{\odot}) = 0.0037$ (Kroupa 2002), and the mean stellar mass $M_{\text{mean}} = 0.4 M_{\odot}$. The total SN energy of a halo is given by

$$E_{\text{SN}}^{\text{tot}} = M_*/(\langle m_* \rangle) \times \xi(m_* > 8 M_{\odot}) \times E_{\text{sn}}. \quad (6)$$

From Table 3 we see that for m09, the amount of SNe energy is not sufficient to create a large core at 3 per cent R_{vir} even if it all couples (via secondary gravitational interactions) to the DM. Even creating 1 per cent R_{vir} core is difficult as it requires that more than 20 per cent of the available energy is coupled to DM, which is unlikely given the indirect connection via the change of gravitational potential and that a large fraction of this energy is in heavily mass loaded winds that move rapidly (Muratov et al. 2015).

However, for m10, a small core is energetically possible and even a 3 per cent R_{vir} core requires less than 10 per cent of the available energy. We see some signs of profile flattening within the inner 1 per cent of R_{vir} , but not a fully developed core. However, for the same halo mass Oñorbe et al. (2015) show that a slightly different SF history can cause a much larger effect and form a central core with a radius of 1–2 per cent R_{vir} .

In m11, the SN energy is three orders of magnitude higher than what is needed to create a small core at 1 per cent of R_{vir} and even 3 per cent of R_{vir} core can be created with few percent of coupling efficiency. It is therefore not surprising that this halo indeed hosts a large core. Even though the depth of the potential well increases in more massive haloes, the amount of stellar mass is a steep function of halo mass, which provides sufficient energy to affect central DM profiles. This is why we see a relatively sharp transition at $\sim 10^{10} M_{\odot}$ from low-mass cuspy haloes to core-like haloes at higher masses.

5.2 Correlation with star formation

Here we show a simple connection between bursts of SF that cause strong gas ejection episodes, and the change in the amount of DM in the central part of haloes. In Fig. 12, we plot the change in the enclosed DM mass as a function of the peak SFR in m11, the halo in which the effect of feedback on DM is one of the strongest in

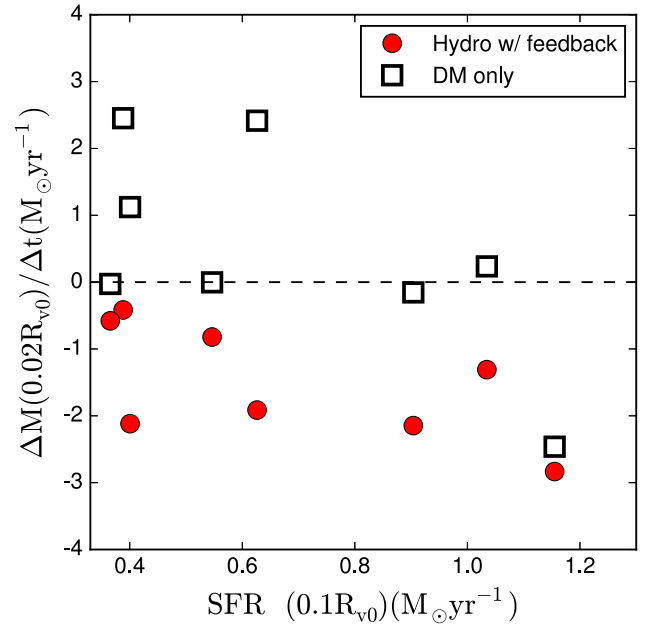


Figure 12. The rate of change of enclosed DM mass within 2 per cent of R_{v0} as a function of peak SFR within 10 per cent R_{v0} in m11 (red circles) and **dam11** (black open squares) simulations, both averaged over 0.2 Gyr. In both cases we plot the SFR from the corresponding feedback simulations at the same cosmic time. The SFR is measured 0.2 Gyr ahead of the rate of change of enclosed DM mass. Strong bursts of star formation are followed by a reduction in the enclosed DM mass in the simulation with feedback.

our sample. We focus on bursts with high peaks of SF and neglect low SFR episodes as they typically do not show strong outflows (Muratov et al. 2015). In Fig. 12 we plot the rate of change in the central amount of DM as a function of the peak SFR. The SFR and mass change are measured as averages over 0.2 Gyr. The rate of change of the DM mass is

$$\frac{\Delta M}{\Delta t}(t) = \frac{m_{\text{dm}}(t + 0.2 \text{ Gyr}) - m_{\text{dm}}(t)}{0.2 \text{ Gyr}}, \quad (7)$$

where $m_{\text{dm}}(t)$ is the enclosed DM mass within 2 per cent R_{v0} (the virial radius at $z = 0$) averaged over 0.2 Gyr and t is the beginning time of the interval. We measure SF between $t - 0.2$ Gyr and t , within a larger radius of $0.1 R_{v0}$ to make sure to include stars that move out of the very centre between the two time intervals.¹²

A significant decrease of mass enclosed in the central 2 per cent R_{v0} occurs just after a strong starburst when a large amount of DM is ‘heated’ and effectively pushed out of the central region. At the corresponding times DM-only simulation does not show negative mass change, except in the case of the largest burst which is triggered by the close passage during a merger (as indicated in Fig. 5), i.e. the merger dynamically alters the profile. Hence we conclude that there is a correlation between mass removal from the centre and SFR.

¹² We have also tried to average SFR over 0.1 Gyr and considered the delay of 0.1 Gyr. The enclosed DM mass drops 0.1 Gyr after the peak SFR in a way similar to Fig. 12 (i.e. it is negative), but with the larger scatter and greater changes in DM mass. We note that this simple measurement cannot use much longer intervals for averaging of SF and for time delay because on large time-scales the SFR bursts would be ‘washed out’ and we could have multiple burst episodes within a long delay time interval.

This correlation suggests that a strong burst of SF can provide sufficient feedback to remove a significant amount of baryons, which then cause a decrease in the central potential and lower the concentration of DM. This scenario is consistent with the mechanism suggested in Pontzen & Governato (2012) in which repeated changes of central gravitational potential transfer energy to the orbits of DM particles causing a central density decrease.

Some fraction of the removed DM does return, i.e. cores get partially rebuilt; however, when large cores are present this is a relatively small effect. When cores are established after the period of rapid early halo growth, they survive largely intact for at least several Gyr, before a new major burst of SF and gas expulsion sets in (see Section 3.4). In more massive haloes (m12), shallow profiles that form at higher redshift steepen by redshift zero significantly. In these haloes we see continuous late-time SF without significant galactic wind episodes (Muratov et al. 2015) which continuously increases central density of baryons that dominate the central potential. As a consequence DM haloes are pulled inwards by adiabatic contraction changing shallow profiles into cusps.

While a burst-driven core formation mechanism is consistent with our findings, we cannot exclude a contribution of other dynamical mechanisms, such as the motion of dense baryonic clumps within galaxies with respect to the halo centres. On the other hand, core formation via enhanced dynamical friction from the dense infalling substructure (e.g. El-Zant et al. 2001) is unlikely to play a significant role for the halo mass range probed here, because feedback lowers the density of infalling subhaloes relative to their DM-only counterparts (see Fig. 6; all of the infalling subhaloes in our sample have $M_h \lesssim 10^{11} M_\odot$).

5.3 Significance for the dark matter detection

The DM profile in the Milky Way is important in studies of indirect detection of DM particles from annihilation or decay signals. Recently, extended emission in gamma rays from the galactic centre has been reported based on the data from the Large Area Telescope aboard the *Fermi Gamma Ray Space Telescope* (e.g. Hooper & Goodenough 2011). To interpret this signal as a consequence of annihilation of DM particles or to constrain its contribution it would be extremely useful to know the central DM profile (Abazajian et al. 2014). The signal is consistent with the annihilation rate of thermally produced weakly interacting massive particle DM (Goodenough & Hooper 2009; Boyarsky, Malyshev & Ruchayskiy 2011; Hooper & Goodenough 2011; Hooper & Linden 2011; Abazajian & Kaplinghat 2012; Gordon & Macías 2013; Abazajian et al. 2014; Macías & Gordon 2014) although other possibilities remain valid alternatives (Abazajian 2011).

While we have a small number of objects at the relevant mass, one robust finding from our simulations is that central density of CDM will not correspond to contracted NFW profiles in Milky Way-mass haloes. This helps constrain the range of values used in the modelling of the observed signatures. We find values of $\alpha \sim -1$ to -1.4 at $1-2$ per cent R_{vir} which is consistent with the best fit in Abazajian et al. (2014). However our results from m12 haloes suggest that deeper in the halo, the DM profile is likely shallower $\alpha \sim -0.7$ to -1.1 at <1 per cent R_{vir} . Given the slight differences of the central slope definition and fitting procedures it would be interesting to test if the profiles shown here provide a good match for the observed signal. We defer such more detailed comparison for future work.

5.4 Limitations

There are some clear limitations in our study of DM halo properties. The number of simulated haloes in our sample is limited because such high-resolution simulations are very time-consuming. In future we plan to simulate a much larger number of haloes to be able to extract direct statistical predictions and to compare to observed values, including the scatter in observations and theoretical predictions.

Further limitation comes from the finite resolution. Our simulations are amongst the highest resolution cosmological simulations at $z = 0$ to date and our m09 and m10 simulations can robustly resolve DM profiles on <200 pc scales. However, detailed comparison to observations requires converged results and the ability to exactly integrate circular orbits over a Hubble time without N -body effects within 300 pc in $M_h \sim 10^{10-11} M_\odot$ ($M_* \sim 10^{7-9} M_\odot$) haloes (Walter et al. 2008). In our m10h1297, m10h1146 and m11 simulations, this is larger than many gravitational softening lengths; however, the more stringent Power convergence criterion (Section 2.4; Power et al. 2003) requires a further factor of several improvements in mass resolution to reach this limit.

Resolution requirements can be even more dramatic if one wants to directly study the inner density profiles of small dwarfs that form within subhaloes of large spiral galaxies (e.g. to directly address TBTF problem). Such study would require simulating $M_h = 10^{12} M_\odot$ haloes with $>10^9$ particles within the virial radius. We have shown, however, that many of the CDM ‘problems’ can be explained by baryonic feedback in isolated, well-resolved, dwarf galaxies suggesting a similar solution for satellite galaxies.

We caution that the Power convergence criterion was derived for different time-stepping algorithm, different force accuracy and different softening values. Based on the DM profiles in Fig. 1 we see that steep central slope in DM-only simulations continue to at least factor of ~ 2 smaller radii than what is estimated by Power convergence radius. More importantly, direct resolution tests confirm the convergence in DM cusp profiles down to a factor $\sim 2-3$ smaller radii for our fiducial resolution. This suggests that this convergence criterion might be too conservative for our simulation setup.

We did not study m13 halo ($M_h \sim 10^{13} M_\odot$) from the FIRE project sample because the SFR of this halo, at low redshift, is higher than the observed rates in observed galaxies of the same mass (H14). Some additional physics, e.g. active galactic nucleus (AGN) feedback, is needed to explain the observations. Yet, the investigation of the profiles of these high-mass haloes is highly intriguing both observationally and theoretically. Observationally, gravitational lensing provides an accurate measurement of the enclosed mass of those haloes (Bolton et al. 2008), though DM only constitutes a minor fraction of the mass.

Nevertheless, stellar kinematics and strong lensing do suggest cores in galaxy clusters with $\alpha \sim -0.5$ (Newman et al. 2013a,b). It is not clear if a single mechanism can explain such shallow slopes over a large range. Mechanisms ranging from more frequent major mergers as well as processes such as AGN feedback, magnetic fields, anisotropic conduction and cosmic rays, which are not yet incorporated in our simulations may be important in regulating late-time SF and affecting the core formation in massive haloes (e.g. Peirani, Kay & Silk 2008; Martizzi, Teyssier & Moore 2013).

AGN feedback in particular might even affect haloes with masses similar to our m12 simulations; e.g. Velliscig et al. (2014) showed that AGN feedback can have a non-negligible impact on the halo properties (i.e. mass and profile) down to $M_h \sim 5 \times 10^{11} M_\odot$. These results suggest that the effect of AGN feedback, in addition

to stellar feedback, could further lower the central density of the most massive haloes in our sample. We note that regardless of the dominant feedback mechanism, the overall efficiency of feedback must be similar to what is seen in our simulations, as this efficiency is constrained by the observed M_* – M_h relation. Our simulations provide a clean test for the effects of stellar feedback alone on the DM distribution.

5.5 Dependence on star formation history

Oñorbe et al. (2015) compared our m10 simulation from H14 with the one with a slightly different SN feedback coupling at smaller scales. In our default case, energy deposition is volume weighted while in the other version it was mass weighted. This creates slight differences in the feedback and changes late-time SF. A 1 kpc core was formed in a halo with more prominent late-time SF, while our default m10 simulation shows a core at <400 pc and much higher central density of DM. It is likely that this strong sensitivity is caused by this halo mass being at the transition region from smaller to larger cores.

This then suggests that when a large sample of simulated haloes is available, comparison to the observations around this mass scale will potentially help distinguish between feedback models by analysing their SF histories and properties of their DM haloes (Fitts et al., in preparation). For slightly more massive haloes we see core formations in all cases we explored, regardless of their detailed SF history and the details of their feedback implementations.

6 CONCLUSIONS

We have explored CDM profiles in simulations with stellar feedback. We used the FIRE suite of hydrodynamic simulations initially discussed in H14 and supplement these with four new dwarf galaxy simulations. We also run collisionless counterparts for all of these simulations. We show that baryonic simulations can successfully produce results consistent with observations and alleviate or solve several so-called ‘problems’ of the CDM: ‘cusp-core’, the lack of adiabatic contraction or ‘halo expansion’ and the ‘TBTF’, without any fine tuning or introduction of adjustable parameters. Our main results are as follows.

(1) The baryons have little influence on haloes with $M_h \ll 10^{10} M_\odot$ because only a small fraction of available baryons are converted to stars, owing to feedback and the UV background that suppress their SF after the reionization (Oñorbe et al. 2015; Wheeler et al. 2015). The smallest haloes are therefore perfect places for testing various theories of DM.

(2) The central slopes of DM density profiles are governed by halo mass and stellar mass. Profiles are shallow with relatively large cores for $M_h \sim 10^{10}$ to few times $10^{11} M_\odot$ and $M_* \sim 10^7$ to few times $10^9 M_\odot$, where $\alpha \sim -0.5$ – 0 and cores are $r_{\text{core}} > 1$ kpc. Small central cores can also form at slightly lower masses $M_* \sim 10^6 M_\odot$. This result is consistent with the observations of dwarf galaxies and can explain the ‘cusp-core’ problem.

(3) Bursts of SF and feedback start forming cores at early times but the cores are established typically at later times, e.g. in our m11 simulation core is still growing at $z < 1$. Stable cores are established once central regions of haloes stop their rapid growth. After this time ($z \lesssim 2$) removal of mass from the central region leaves a long-term effect on the halo profile. We show that strong bursts of SF are correlated with DM expansion.

(4) The total SN energy in haloes with $M_h > 10^{10} M_\odot$ is sufficient to produce a core with radius 3 per cent R_{vir} , but not sufficient to make large cores in lower mass haloes. In practice only a few per cent of the available energy is transferred into evacuation of DM from the central region.

(5) Baryonic contraction of DM haloes becomes significant when central regions of haloes are clearly dominated by baryons, which in our simulations occurs for $M_h > 5 \times 10^{11} M_\odot$. However feedback in the progenitors of these massive galaxies significantly lowered central DM density at $z \sim 1$ – 1.5 . The cumulative effect of feedback and contraction is then a profile that in our m12 runs is slightly shallower than NFW. This explains why the normalization of the TFR requires no contraction or halo expansion with respect to the collisionless NFW profile.

(6) Stellar feedback in galaxies with $M_* \sim$ few times 10^6 – $10^8 M_\odot$ lowers the central density of DM when compared to DM-only simulations and significantly reduces the rotational velocity near the centre. This means that relatively low circular velocities of observed galaxies should correspond to much higher maximum circular velocities or virial velocities of haloes or subhaloes in collisionless CDM simulation. This can solve or at least substantially alleviate the ‘TBTF’ problem.

ACKNOWLEDGEMENTS

We would like to thank N. Murray, J. Bullock, M. Boylan-Kolchin and P. Salucci for useful discussion. DK and TKC were supported in part by NSF grant AST-1412153, Hellman Fellowship and funds from the University of California San Diego. Support for PFH was provided by an Alfred P. Sloan Research Fellowship, NASA ATP Grant NNX14AH35G, and NSF Collaborative Research Grant no. 1411920. EQ was supported in part by NASA ATP grant 12-APT12-0183, a Simons Investigator award from the Simons Foundation, and the David and Lucile Packard Foundation. CAFG was supported by NSF through grant AST-1412836, by NASA through grant NNX15AB22G, and by Northwestern University funds. The simulation presented here used computational resources granted by the Extreme Science and Engineering Discovery Environment (XSEDE), which is supported by National Science Foundation grant no. OCI-1053575, specifically allocations TG-AST120025 (PI Kereš), TG-AST130039 (PI Hopkins) and TG-AST1140023 (PI Faucher-Giguère).

REFERENCES

- Abazajian K., 2006, *Phys. Rev. D*, 73, 063513
- Abazajian K. N., 2011, *J. Cosmol. Astropart. Phys.*, 3, 10
- Abazajian K. N., Kaplinghat M., 2012, *Phys. Rev. D*, 86, 083511
- Abazajian K. N., Canac N., Horiuchi S., Kaplinghat M., 2014, *Phys. Rev. D*, 90, 023526
- Agertz O. et al., 2007, *MNRAS*, 380, 963
- Agertz O., Teyssier R., Moore B., 2009, *MNRAS*, 397, L64
- Barnes J. E., 2012, *MNRAS*, 425, 1104
- Begeman K. G., 1987, PhD thesis, Kapteyn Institute, 1987
- Blumenthal G. R., Faber S. M., Flores R., Primack J. R., 1986, *ApJ*, 301, 27
- Bolton A. S., Burles S., Koopmans L. V. E., Treu T., Gavazzi R., Moustakas L. A., Wayth R., Schlegel D. J., 2008, *ApJ*, 682, 964
- Boyarisky A., Malyshev D., Ruchayskiy O., 2011, *Phys. Lett. B*, 705, 165
- Boylan-Kolchin M., Bullock J. S., Kaplinghat M., 2011, *MNRAS*, 415, L40
- Broeils A. H., 1992, PhD thesis, Univ. Groningen, the Netherlands
- Brook C. B., Di Cintio A., 2015, *MNRAS*, 450, 3920
- Brook C. B., Stinson G., Gibson B. K., Roškar R., Wadsley J., Quinn T., 2012, *MNRAS*, 419, 771

- Brooks A. M., Zolotov A., 2014, *ApJ*, 786, 87
- Bryan G. L., Norman M. L., 1998, *ApJ*, 495, 80
- Burkert A., 2000, *ApJ*, 534, L143
- Carlson E. D., Machacek M. E., Hall L. J., 1992, *ApJ*, 398, 43
- Courteau S., Dutton A. A., 2015, *ApJ*, 801, L20
- Courteau S., Dutton A. A., van den Bosch F. C., MacArthur L. A., Dekel A., McIntosh D. H., Dale D. A., 2007, *ApJ*, 671, 203
- Cullen L., Dehnen W., 2010, *MNRAS*, 408, 669
- Davé R., Spergel D. N., Steinhardt P. J., Wandelt B. D., 2001, *ApJ*, 547, 574
- de Blok W. J. G., McGaugh S. S., 1997, *MNRAS*, 290, 533
- de Blok W. J. G., Walter F., Brinks E., Trachternach C., Oh S.-H., Kennicutt R. C., Jr, 2008, *AJ*, 136, 2648
- de Laix A. A., Scherrer R. J., Schaefer R. K., 1995, *ApJ*, 452, 495
- Dekel A., Devor J., Hetzroni G., 2003, *MNRAS*, 341, 326
- Del Popolo A., 2009, *ApJ*, 698, 2093
- Di Cintio A., Brook C. B., Macciò A. V., Stinson G. S., Knebe A., Dutton A. A., Wadsley J., 2014, *MNRAS*, 437, 415
- Dunstan R. M., Abazajian K. N., Polisensky E., Ricotti M., 2011, preprint ([arXiv:1109.6291](https://arxiv.org/abs/1109.6291))
- Durier F., Dalla Vecchia C., 2012, *MNRAS*, 419, 465
- Dutton A. A., Macciò A. V., 2014, *MNRAS*, 441, 3359
- Dutton A. A., van den Bosch F. C., Dekel A., Courteau S., 2007, *ApJ*, 654, 27
- Dutton A. A., Conroy C., van den Bosch F. C., Prada F., More S., 2010, *MNRAS*, 407, 2
- El-Zant A., Shlosman I., Hoffman Y., 2001, *ApJ*, 560, 636
- Elbert O. D., Bullock J. S., Garrison-Kimmel S., Rocha M., Oñorbe J., Peter A. H. G., 2015, *MNRAS*, 453, 29
- Faucher-Giguère C.-A., Kereš D., 2011, *MNRAS*, 412, L118
- Faucher-Giguère C.-A., Lidz A., Zaldarriaga M., Hernquist L., 2009, *ApJ*, 703, 1416
- Faucher-Giguère C., Kereš D., Dijkstra M., Hernquist L., Zaldarriaga M., 2010, *ApJ*, 725, 633
- Faucher-Giguère C.-A., Hopkins P. F., Keres D., Muratov A. L., Quataert E., Murray N., 2015, *MNRAS*, 449, 987
- Ferrero I., Abadi M. G., Navarro J. F., Sales L. V., Gurovich S., 2012, *MNRAS*, 425, 2817
- Flores R. A., Primack J. R., 1994, *ApJ*, 427, L1
- Garrison-Kimmel S., Rocha M., Boylan-Kolchin M., Bullock J. S., Lally J., 2013, *MNRAS*, 433, 3539
- Garrison-Kimmel S., Boylan-Kolchin M., Bullock J. S., Kirby E. N., 2014, *MNRAS*, 444, 222
- Gentile G., Salucci P., Klein U., Vergani D., Kalberla P., 2004, *MNRAS*, 351, 903
- Gnedin O. Y., Ostriker J. P., 2001, *ApJ*, 561, 61
- Gnedin O. Y., Zhao H., 2002, *MNRAS*, 333, 299
- Gnedin O. Y., Kravtsov A. V., Klypin A. A., Nagai D., 2004, *ApJ*, 616, 16
- Goodenough L., Hooper D., 2009, preprint ([arXiv:0910.2998](https://arxiv.org/abs/0910.2998))
- Gordon C., Macías O., 2013, *Phys. Rev. D*, 88, 083521
- Governato F. et al., 2010, *Nature*, 463, 203
- Governato F. et al., 2012, *MNRAS*, 422, 1231
- Hahn O., Abel T., 2011, *MNRAS*, 415, 2101
- Hooper D., Goodenough L., 2011, *Phys. Lett. B*, 697, 412
- Hooper D., Linden T., 2011, *Phys. Rev. D*, 84, 123005
- Hopkins P. F., 2013, *MNRAS*, 428, 2840
- Hopkins P. F., 2015, *MNRAS*, 450, 53
- Hopkins P. F., Kereš D., Oñorbe J., Faucher-Giguère C.-A., Quataert E., Murray N., Bullock J. S., 2014, *MNRAS*, 445, 581 (H14)
- Hunter D. A. et al., 2012, *AJ*, 144, 134
- Kaplinghat M., Keeley R. E., Linden T., Yu H.-B., 2014, *Phys. Rev. Lett.*, 113, 021302
- Katz N., White S. D. M., 1993, *ApJ*, 412, 455
- Kaufmann T., Mayer L., Wadsley J., Stadel J., Moore B., 2006, *MNRAS*, 370, 1612
- Kereš D., Hernquist L., 2009, *ApJ*, 700, L1
- Kereš D., Katz N., Weinberg D. H., Davé R., 2005, *MNRAS*, 363, 2
- Kereš D., Vogelsberger M., Sijacki D., Springel V., Hernquist L., 2012, *MNRAS*, 425, 2027
- Kim J.-h. et al., 2014, *ApJS*, 210, 14
- Kirby E. N., Bullock J. S., Boylan-Kolchin M., Kaplinghat M., Cohen J. G., 2014, *MNRAS*, 439, 1015
- Knebe A., Green A., Binney J., 2001, *MNRAS*, 325, 845
- Knollmann S. R., Knebe A., 2009, *ApJS*, 182, 608
- Kochanek C. S., White M., 2000, *ApJ*, 543, 514
- Kroupa P., 2002, *Science*, 295, 82
- Krumholz M. R., Gnedin N. Y., 2011, *ApJ*, 729, 36
- Laporte C. F. P., Peñarrubia J., 2015, *MNRAS*, 449, L90
- Leitherer C. et al., 1999, *ApJS*, 123, 3
- Loeb A., Weiner N., 2011, *Phys. Rev. Lett.*, 106, 171302
- Lovell M. R. et al., 2012, *MNRAS*, 420, 2318
- Macciò A. V., Paduroiu S., Anderhalden D., Schneider A., Moore B., 2012a, *MNRAS*, 424, 1105
- Macciò A. V., Stinson G., Brook C. B., Wadsley J., Couchman H. M. P., Shen S., Gibson B. K., Quinn T., 2012b, *ApJ*, 744, L9
- Machacek M. E., Carlson E. D., Hall L. J., 1993, *Ann. New York Acad. Sci.*, 688, 681
- Macías O., Gordon C., 2014, *Phys. Rev. D*, 89, 063515
- Madau P., Shen S., Governato F., 2014, *ApJ*, 789, L17
- Martizzi D., Teyssier R., Moore B., 2013, *MNRAS*, 432, 1947
- Martizzi D., Faucher-Giguère C.-A., Quataert E., 2015, *MNRAS*, 450, 504
- Mashchenko S., Couchman H. M. P., Wadsley J., 2006, *Nature*, 442, 539
- Miralda-Escudé J., 2002, *ApJ*, 564, 60
- Moore B., 1994, *Nature*, 370, 629
- Muratov A. L., Keres D., Faucher-Giguère C.-A., Hopkins P. F., Quataert E., Murray N., 2015, preprint ([arXiv:1501.03155](https://arxiv.org/abs/1501.03155))
- Navarro J. F., Eke V. R., Frenk C. S., 1996, *MNRAS*, 283, L72
- Navarro J. F., Frenk C. S., White S. D. M., 1997, *ApJ*, 490, 493
- Nelson D., Vogelsberger M., Genel S., Sijacki D., Kereš D., Springel V., Hernquist L., 2013, *MNRAS*, 429, 3353
- Newman A. B., Treu T., Ellis R. S., Sand D. J., 2013a, *ApJ*, 765, 25
- Newman A. B., Treu T., Ellis R. S., Sand D. J., Nipoti C., Richard J., Jullo E., 2013b, *ApJ*, 765, 24
- Ogiya G., Mori M., 2011, *ApJ*, 736, L2
- Ogiya G., Mori M., 2014, *ApJ*, 793, 46
- Oh S.-H., de Blok W. J. G., Brinks E., Walter F., Kennicutt R. C., Jr, 2011, *AJ*, 141, 193
- Oh S.-H. et al., 2015, *AJ*, 149, 180
- Oñorbe J., Garrison-Kimmel S., Maller A. H., Bullock J. S., Rocha M., Hahn O., 2014, *MNRAS*, 437, 1894
- Oñorbe J., Boylan-Kolchin M., Bullock J. S., Hopkins P. F., Kereš D., Faucher-Giguère C.-A., Quataert E., Murray N., 2015, preprint ([arXiv:1502.02036](https://arxiv.org/abs/1502.02036))
- Oppenheimer B. D., Davé R., Kereš D., Fardal M., Katz N., Kollmeier J. A., Weinberg D. H., 2010, *MNRAS*, 406, 2325
- Papastergis E., Giovanelli R., Haynes M. P., Shankar F., 2015, *A&A*, 574, A113
- Peirani S., Kay S., Silk J., 2008, *A&A*, 479, 123
- Peñarrubia J., Pontzen A., Walker M. G., Koposov S. E., 2012, *ApJ*, 759, L42
- Peter A. H. G., Rocha M., Bullock J. S., Kaplinghat M., 2013, *MNRAS*, 430, 105
- Pontzen A., Governato F., 2012, *MNRAS*, 421, 3464
- Pontzen A., Governato F., 2014, *Nature*, 506, 171
- Porter D. H., 1985, PhD thesis, California Univ., Berkeley
- Power C., Navarro J. F., Jenkins A., Frenk C. S., White S. D. M., Springel V., Stadel J., Quinn T., 2003, *MNRAS*, 338, 14
- Price D. J., 2008, *J. Comput. Phys.*, 227, 10040
- Price D. J., Monaghan J. J., 2007, *MNRAS*, 374, 1347
- Rahmati A., Pawlik A. H., Raicevic M., Schaye J., 2013, *MNRAS*, 430, 2427
- Read J. I., Gilmore G., 2005, *MNRAS*, 356, 107
- Rocha M., Peter A. H. G., Bullock J. S., Kaplinghat M., Garrison-Kimmel S., Oñorbe J., Moustakas L. A., 2013, *MNRAS*, 430, 81
- Romano-Díaz E., Shlosman I., Hoffman Y., Heller C., 2008, *ApJ*, 685, L105
- Saitoh T. R., Makino J., 2013, *ApJ*, 768, 44
- Salucci P., Burkert A., 2000, *ApJ*, 537, L9

- Sijacki D., Vogelsberger M., Kereš D., Springel V., Hernquist L., 2012, MNRAS, 424, 2999
- Sommer-Larsen J., 2006, ApJ, 644, L1
- Spekkens K., Giovanelli R., Haynes M. P., 2005, AJ, 129, 2119
- Spergel D. N. et al., 2007, ApJS, 170, 377
- Springel V., 2005, MNRAS, 364, 1105
- Springel V. et al., 2005, Nature, 435, 629
- Steinmetz M., Navarro J., 2000, in Combes F., Mamon G. A., Charmandaris V., eds, ASP Conf. Ser. Vol. 197, Dynamics of Galaxies: From the Early Universe to the Present. Astron. Soc. Pac., San Francisco, p. 165
- Strigari L. E., Kaplinghat M., Bullock J. S., 2007, Phys. Rev. D, 75, 061303
- Swaters R. A., Madore B. F., van den Bosch F. C., Balcells M., 2003, ApJ, 583, 732
- Teyssier R., Pontzen A., Dubois Y., Read J. I., 2013, MNRAS, 429, 3068
- Tollet E. et al., 2015, preprint (arXiv:1507.03590)
- Tonini C., Lapi A., Salucci P., 2006, ApJ, 649, 591
- Tully R. B., Fisher J. R., 1976, BAAS, 8, 555
- Velliscig M., van Daalen M. P., Schaye J., McCarthy I. G., Cacciato M., Le Brun A. M. C., Dalla Vecchia C., 2014, MNRAS, 442, 2641
- Verheijen M. A. W., 1997, PhD thesis, Univ. Groningen, the Netherlands
- Vogelsberger M., Zavala J., Loeb A., 2012a, MNRAS, 423, 3740
- Vogelsberger M., Sijacki D., Kereš D., Springel V., Hernquist L., 2012b, MNRAS, 425, 3024
- Walker M. G., Mateo M., Olszewski E. W., Peñarrubia J., Wyn Evans N., Gilmore G., 2009, ApJ, 704, 1274
- Walter F., Brinks E., de Blok W. J. G., Bigiel F., Kennicutt R. C., Jr, Thornley M. D., Leroy A., 2008, AJ, 136, 2563
- Weisz D. R., Dolphin A. E., Skillman E. D., Holtzman J., Gilbert K. M., Dalcanton J. J., Williams B. F., 2014, ApJ, 789, 147
- Wheeler C., Onorbe J., Bullock J. S., Boylan-Kolchin M., Elbert O. D., Garrison-Kimmel S., Hopkins P. F., Keres D., 2015, MNRAS, 453, 1305
- Wolf J., Martinez G. D., Bullock J. S., Kaplinghat M., Geha M., Muñoz R., Simon J. D., Avedo F. F., 2010, MNRAS, 406, 1220
- Yoshida N., Springel V., White S. D. M., Tormen G., 2000, ApJ, 535, L103
- Zolotov A. et al., 2012, ApJ, 761, 71

APPENDIX A: CONVERGENCE LIMITS

We use the Power convergence criterion (Power et al. 2003) to derive empirical formulae for the minimum particle mass needed to quantify cusps down to 0.2–3 per cent R_{vir} . We consider DM-only simulations and assume their profiles can be fitted with an NFW profile. The ratio between the scale radius r_s and the virial radius R_{vir} is determined through a concentration–mass relation from Dutton & Macciò (2014), and we use the virial overdensity from Bryan & Norman (1998). Then we calculate the enclosed number of particles and density within 0.3–3 per cent R_{vir} as well as 300 and 700 pc. From equation (1), we estimate the minimum radius such that $t_{\text{relax}}/t_0 < 0.6$ and plot the relation between the required particle mass to meet this criteria at the desired radius, and the total halo mass in Fig. A1.

The data can be fitted with a linear relation,

$$\log_{10}(m_{\text{dm}}/M_{\odot}) = a \log_{10}(M_{\text{h}}/M_{\odot}) + b, \quad (\text{A1})$$

where a and b are listed in Table A1.

Oh et al. (2011) measured α between 300–700 pc in dwarf galaxies with M_* between 10^6 – $10^9 M_{\odot}$ and found core profiles in those galaxies. In order to match this observational result, the minimum resolved radius in the simulations should be around 300 pc. For a fixed physical radius this turns into a requirement in particle mass that is almost independent of the halo mass, owing to higher concentration in lower mass haloes. The maximum mass of DM particles needs to be slightly smaller than $10^4 M_{\odot}$ to converge at 300 pc and

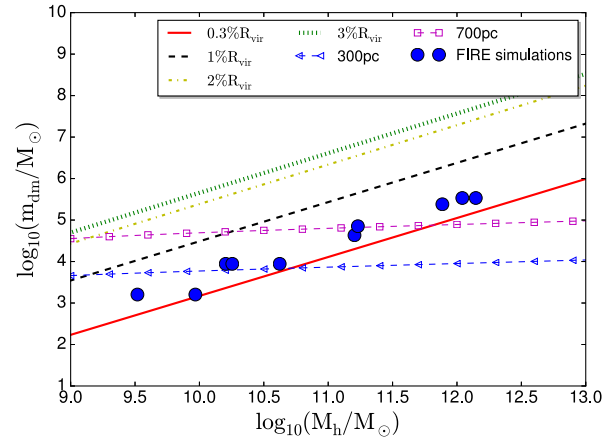


Figure A1. The virial mass of a halo as a function of the maximum particle mass allowed in order to reach convergent profiles at 0.3, 1, 2, 3 per cent R_{vir} , 300 and 700 pc. Small circles represent the corresponding values from DM-only simulation listed in Table 1. (Our full physics simulations have the same DM particle numbers.)

Table A1. The coefficients in equation (A1) for different resolutions.

$r_{\text{res}}/R_{\text{vir}}$	a	b
0.3 per cent	0.94	−6.2
1 per cent	0.95	−5.0
2 per cent	0.95	−4.3
3 per cent	0.96	−3.9

smaller than $10^5 M_{\odot}$ to converge within 700 pc. Our m09 and m10 are clearly converged at both of these radii.

Our slightly more massive haloes m10h1297, m10h1146 and m10h573 are marginally converged at 300 pc, but fully converged at 700 pc. Fig. A1 assumes an NFW profile and implicitly assumes that the central region has close to a Hubble time to undergo relaxation processes. As discussed in the main text, it is not clear what the appropriate convergence criterion should be once large cores are formed and the central density is reduced. This likely depends on the core formation time as well as details of the gravitational softening of multiple particle species and their time-stepping algorithms.

It is important to note that our DM force softening is typically a factor of five smaller than the corresponding Power convergence criterion. Furthermore, we have also tested if the force softening of the baryonic component influences DM profiles: e.g. we increased the baryonic softening from 2.0 to 25 pc in the slightly modified version of m10 run in Onorbe et al. (2015) and found that the DM profile was only mildly changed (the core size was actually larger in run with smaller softening). While two-body relaxation effects are important in estimating central DM profiles, in H14 (appendix C) we have used idealized runs to show that our standard resolution in m12 runs is also sufficient to reliably determine other relevant quantities, e.g. SFR, wind mass-loading and gas phase distribution. All are consistent to within a factor of ~ 2 even with ~ 50 times better mass resolution. This indicates that the general properties of DM profiles on resolved scales in our simulations are numerically robust.

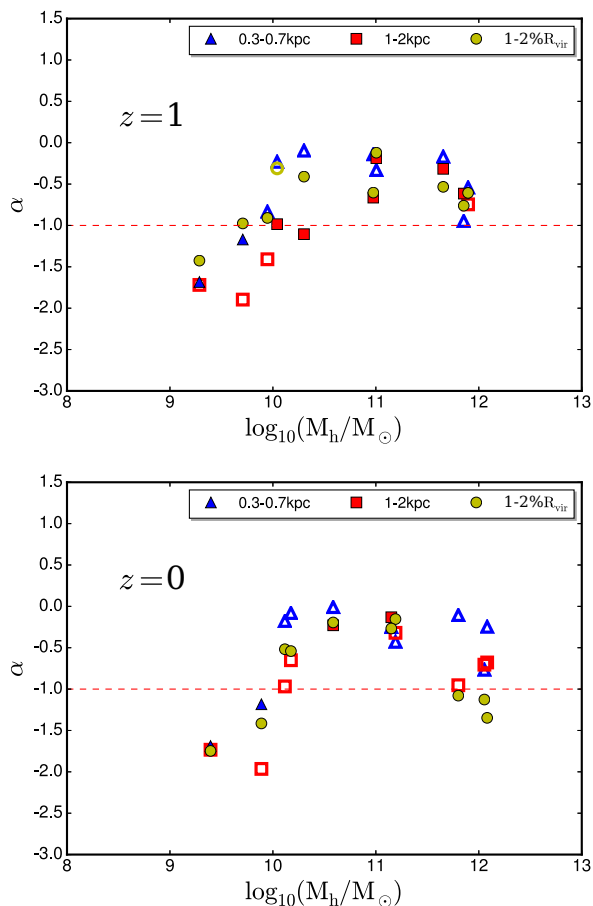


Figure B1. DM profile slope α inferred from different criteria plotted as a function of the halo mass. ‘0.3–0.7 kpc’ is the DM density slope interpolated within 0.3–0.7 kpc from the centre. ‘1–2 per cent R_{vir} ’ is α interpolated within 1–2 per cent R_{vir} . ‘1–2 kpc’ is α interpolated within 1–2 kpc. Filled circles show that the profile measurement range is larger than $0.5r_{\text{pow}}$ and smaller than a third of r_s . Open circles indicate that one of these criteria is not satisfied.

APPENDIX B: CHOICE OF α

We investigate the effect of different fitting ranges on α in this appendix. We consider three different fitting ranges, 1–2 kpc, 0.3–0.7 kpc and 1–2 per cent R_{vir} . Fig. B1 shows α that corresponds to those ranges. In general, 0.3–0.7 kpc includes some overlap

below the Power radius for haloes with mass larger than $10^{11} M_{\odot}$ but α in this range can be directly compared with observations of dwarf galaxies. 1–2 kpc lies outside the central region ($>1.3r_s$) in small dwarfs. 1–2 per cent R_{vir} is well resolved, lies inside the central region and is physically meaningful, so we use this fitting range as our ‘default’ choice in the main text. Overall, all of the methods show very similar trends of the DM density profile slope with mass.

This paper has been typeset from a \LaTeX file prepared by the author.